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# AKARI/IRC NEAR-INFRARED SPECTRAL ATLAS OF GALACTIC PLANETARY NEBULAE

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## ABSTRACT

Near-infrared (2.5–5.0  $\mu\text{m}$ ) low-resolution ( $\lambda/\Delta\lambda \sim 100$ ) spectra of 72 Galactic planetary nebulae (PNe) were obtained with the Infrared Camera (IRC) in the post-helium phase. The IRC, equipped with a  $1' \times 1'$  window for spectroscopy of a point source, was capable of obtaining near-infrared spectra in a slit-less mode without any flux loss due to a slit. The spectra show emission features including hydrogen recombination lines and the 3.3–3.5  $\mu\text{m}$  hydrocarbon features. The intensity and equivalent width of the emission features were measured by spectral fitting. We made a catalog providing unique information on the investigation of the near-infrared emission of PNe. In this paper, details of the observations and characteristics of the catalog are described.

**Key words:** catalogs – planetary nebulae: general

## 1. INTRODUCTION

Planetary nebulae (PNe) are a late evolutionary stage of low- and intermediate-mass stars (e.g., Blöcker 1995; Schönberner & Blöcker 1996). They are surrounded by rich circumstellar material (CSM) ejected during the Asymptotic Giant Branch (AGB) phase (Schönberner et al. 2005). The CSM is eventually incorporated into the interstellar medium and consists of the ingredients for next-generation star-forming activity.

Infrared spectra are rich in CSM emission features. These features can be used to investigate the mass, temperature, and chemical properties of CSMs (e.g., Pottasch et al. 1984; Bernard-Salas & Tielens 2005; Pottasch & Bernard-Salas 2006; Phillips & Márquez-Lugo 2011). In some PNe, far-infrared emission is dominated by that from dust grains, providing information about the temperature and the total mass of dust grains. Spectroscopic observations in the mid-infrared ( $\sim 5\text{--}40\ \mu\text{m}$ ) have been used to identify dust species (e.g., Volk & Cohen 1990; Stanghellini et al. 2012). Emission from stochastically heated dust grains also appears in the mid-infrared (Dwek et al. 1997; Draine & Li 2001). The dust features commonly found are PAHs, MgS, silicate, and some potential cases of SiC. Near-infrared ( $\sim 1\text{--}5\ \mu\text{m}$ ) continuum emission in PNe involves free-free emission, stellar continuum, and hot dust emission. Hydrogen recombination lines (RLs) are prominent in the near-infrared, such as Brackett- $\alpha$  at 4.051  $\mu\text{m}$ . The PAH emission feature appears at 3.3  $\mu\text{m}$  (Tokunaga et al. 1991; Beintema et al. 1996; Roche et al. 1996), while there is also emission from aliphatic C–H bonds at around 3.4–3.5  $\mu\text{m}$  (Joblin et al. 1996; Sloan et al. 1997).

Infrared space telescopes have been instrumental for PNe observations despite limited spatial resolution. Eliminating the effects of heavy terrestrial atmospheric absorption provides the critical advantage of continuous infrared spectroscopic coverage. Volk & Cohen (1990) investigated 7–23  $\mu\text{m}$  low-resolution spectra of 170 Galactic PNe with the *Infrared Astronomical Satellite* (IRAS). The Short-Wave Spectrometer (SWS) and the Long-Wave Spectrometer (LWS) on board the *Infrared Space Observatory* (ISO) cover 2.5–45 and 45–197  $\mu\text{m}$ . A number of emission lines have been identified and the physical conditions of circumstellar nebulae have been

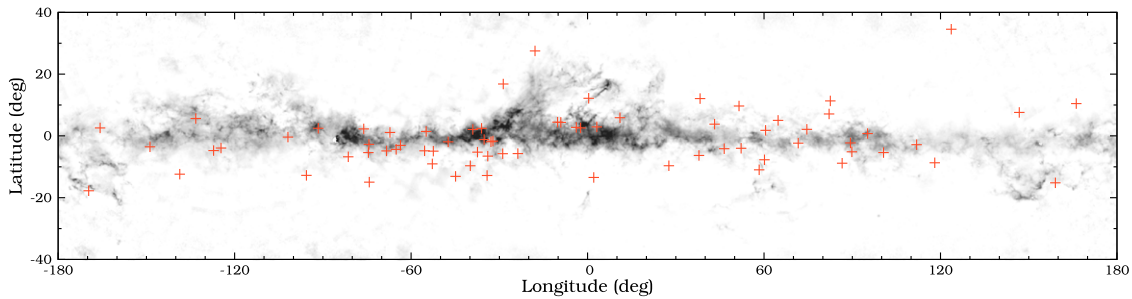
investigated (e.g., Bernard-Salas & Tielens 2005; Pottasch & Bernard-Salas 2006). The Infrared Spectrograph (IRS) on board the *Spitzer Space Telescope* covers 5.5–37  $\mu\text{m}$ . Stanghellini et al. (2012) obtained 157 spectra of compact Galactic PNe and investigated their dust emission. Spectra of PNe in the Magellanic Clouds were also obtained with the IRS (Bernard-Salas et al. 2008, 2009; Stanghellini & Haywood 2010) thanks to its high sensitivity. The ISO/SWS had obtained 2.5–5.0  $\mu\text{m}$  spectra of Galactic PNe. About 85 spectra of Galactic PNe are available in Sloan et al. (2003). However, only a few PNe have spectra with a sufficient signal-to-noise ratio (S/N) at 2.5–5.0  $\mu\text{m}$  to investigate weak features such as the 3.4–3.5  $\mu\text{m}$  aliphatic emission. The Infrared Camera (IRC) on board the AKARI satellite, for the first time, provides near-infrared (2.5–5.0  $\mu\text{m}$ ) spectroscopy with a sensitivity of a few mJy (Onaka et al. 2007).

The present paper reports near-infrared spectroscopy with the AKARI/IRC. Observations were carried out as part of an Open Time Program for the post-helium phase, “Near-Infrared Spectroscopy of Planetary Nebulae” (PNSPC). Near-infrared spectra of 72 PNe were obtained. The data were compiled as a near-infrared spectral catalog. In this paper, we describe details of the observations and the characteristics of the catalog. Detailed information about the observations is given in Section 2. Section 3 describes the reduction procedures and the method to extract the spectra. The characteristics of the retrieved spectra and the format of the compiled catalog are presented in Section 4. The quality of the IRC spectra and the statistical properties of the catalog are discussed in comparison with the observations by other facilities in Section 5. We summarize the paper in Section 6.

## 2. OBSERVATIONS

### 2.1. Target Selection

The observed targets were selected from the Strasbourg-ESO Catalog of Galactic PNe (Acker et al. 1992), taking into account their NIR magnitude and angular size. The spectroscopic sensitivity of the IRC is a few mJy in the range of 2–5  $\mu\text{m}$ , while the saturation limit is about 10 Jy. With these limits taken into account, objects whose *K*-band magnitude is



**Figure 1.** Distribution of the observed planetary nebulae (the crosses) overlaid on an extinction map in the Galactic coordinates created from the DSS (Dobashi et al. 2005a). The gray scale indicates extinction at the V-band from 0.1 to 5 mag.

between 5 and 13 were selected as targets for this program. The observations were carried out with the  $1' \times 1'$  Np-window (see Section 2.2 and Onaka et al. 2007) in the slit-less mode. Therefore, compact objects were preferred to avoid degrading the spectral resolution. The candidate list was then narrowed to objects whose radius in the optical was smaller than  $5''$ , comparable to the full width half maximum (FWHM) of the IRC point-spread function (PSF) ( $\sim 3''$ ), to exclude apparently large objects. Finally, 72 objects were selected and their spectra were successfully obtained. Figure 1 shows the distribution of the targets in Galactic coordinates. The PNG ID, the AKARI observation ID, and the coordinates of the PNSPC samples are listed in Table 1. Table 2 summarizes miscellaneous information of the objects, including the optical (the V-band) and infrared (the 2MASS  $K_s$ -band) magnitudes.

The target selection method was independent of the chemistry and dust compositions of the circumstellar nebula. The catalog is not considered to be biased, in terms of chemical abundance and the chemistry of dust. However, the objects were selected based on their sizes. Since PNe expand as they evolve, the selection could be biased toward young PNe. The potential biases of the PNSPC samples are discussed in detail in Section 5.

## 2.2. Instruments

Although the AKARI/IRC has three channels (Onaka et al. 2007), only the NIR channel was available during the post-helium phase. The field of view of the NIR channel consists of four parts: the N/Nc-window, the Ns-slit, the Np-window, and the Nh-slit (see Figure 3 in Onaka et al. 2007). The observations were performed with the Np-window, which was designed for point source, slit-less spectroscopy. Spectroscopy with the Np-window allows us to collect all of the flux from an object. The present spectroscopy was carried out with the grism, providing  $2.5\text{--}5.0\ \mu\text{m}$  spectra with the spectral resolution of  $\lambda/\Delta\lambda \sim 100$ , for point sources ( $\lesssim 3''$ ). The spectral resolution was degraded for extended objects.

The observations were carried out with the observation template IRCZ4. This mode consists of three sequences: the first set is four spectroscopic images, the second set is a broadband image, and the final set is another sequence of four spectroscopic images. Thus, each pointing observation involves eight frames for spectroscopy and one frame for broadband imaging. Every frame consists of short (4.58 s) and long (44.41 s) exposure images. Due to the short integration time, the typical S/N of the short exposure images is worse than that of the long exposure images. In general, we used only the long exposure images. The short exposure images were used only

for cases when the long exposure image showed saturated pixels.

Each object was observed either once or twice. Thus, there were either eight or sixteen frames for each object. However, some frames were not usable due to cosmic rays or unstable pointing during the exposure. The net integration time thus depends on the target, listed in the sixth column of Table 1.

## 3. DATA REDUCTION

### 3.1. Bias Subtraction and Flux Calibration

The basic data reduction procedure used here includes subtraction of dark frames, linearity correction, saturation masking, flat fielding, and subtraction of background emission. All steps were carried out with the IRC Spectroscopy Toolkit for Phase 3 (Version 20111121).<sup>5</sup> Pixels affected by cosmic rays were masked by hand. Some pixels were saturated by bright objects in the long exposure images. The saturated pixels were replaced by the corresponding pixels from the short exposure images. Objects whose spectra were saturated in the long exposure images are marked in the eighth column of Table 1. Spectroscopic images were aligned by maximizing the correlation among the images. Then, the images were combined by median stacking, producing a two-dimensional spectrum. Frames that were affected by unstable pointing during the exposure were excluded from stacking. The toolkit also produced a two-dimensional noise map, which was given as a standard deviation of the stacked frames.

### 3.2. Extraction of One-dimensional Spectrum

Some objects were extended with respect to the PSF of the IRC. To make a homogeneous data reduction, we assumed that all the sources were extended. We used a sufficiently wide aperture ( $>10$  pixels) to extract a one-dimensional spectrum that included all of the fluxes from the object and did not apply any aperture correction. The standard error in the spectrum was calculated using the noise map produced by the toolkit.

It was difficult to use a wide aperture for objects which were located in crowded regions. Noble et al. (2013) invented a novel method to extract the spectrum of a certain object in a crowded region by profile fitting. We used a similar method to remove the contribution from neighboring objects. Figure 2 shows a schematic view of the method (hereafter the spectral decomposition procedure). Panel (A) of Figure 2(a) shows an example two-dimensional spectrum of a target in a crowded

<sup>5</sup> Available at <http://www.ir.isas.jaxa.jp/ASTRO-F/Observation/DataReduction/IRC/>.



**Table 1**  
Summary of Observations

PN G	Obs. ID	R.A. (J2000) (hh:mm:ss)	Decl. (J2000) (dd:mm:ss)	$N^a$	$t_{\text{int}}^b$ (s)	FWHM ( $\mu\text{m}$ )	Flags		
							$\mathcal{S}^c$	$\mathcal{D}^d$	$\mathcal{C}^e$
000.3+12.2	3460107	17:01:33.6	−21:49:33.5	17	754.97	0.045	N	N	N
002.0−13.4	3460003	18:45:50.7	−33:20:35.0	16	710.56	0.040	N	Y	N
003.1+02.9	3460005	17:41:52.8	−24:42:07.7	9	399.69	0.045	N	Y	Y
011.0+05.8	3460012	17:48:19.8	−16:28:44.0	18	799.38	0.035	N	N	N
027.6−09.6	3460016	19:16:28.3	−09:02:37.0	9	399.69	0.040	N	Y	N
037.8−06.3	3460019	19:22:56.9	+01:30:48.0	17	754.97	0.035	N	Y	N
038.2+12.0	3460020	18:17:34.0	+10:09:05.0	9	399.69	0.040	N	N	N
043.1+03.8	3460093	18:56:33.6	+10:52:12.0	9	399.69	0.035	N	Y	Y
046.4−04.1	3460021	19:31:16.4	+10:03:21.7	9	399.69	0.040	N	Y	Y
051.4+09.6	3460022	18:49:47.6	+20:50:39.5	17	754.97	0.040	N	N	N
052.2−04.0	3460094	19:42:18.7	+15:09:09.0	18	799.38	0.035	N	Y	N
058.3−10.9	3460023	20:20:08.8	+16:43:54.0	18	799.38	0.040	N	N	N
060.1−07.7	3460024	20:12:42.9	+19:59:23.0	14	621.74	0.048	N	Y	N
060.5+01.8	3460025	19:38:08.4	+25:15:42.0	18	799.38	0.035	N	Y	N
064.7+05.0	3460026	19:34:45.3	+30:30:59.2	17	77.86	0.048	Y	N	N
071.6−02.3	3460027	20:21:03.8	+32:29:24.0	18	799.38	0.035	N	N	N
074.5+02.1	3460028	20:10:52.5	+37:24:41.0	18	799.38	0.040	N	N	N
082.1+07.0	3460029	20:10:23.7	+46:27:39.0	18	799.38	0.045	N	Y	N
082.5+11.3	3460030	19:49:46.6	+48:57:40.0	9	399.69	0.035	N	Y	N
086.5−08.8	3460114	21:33:08.3	+39:38:09.7	18	799.38	0.045	N	Y	N
089.3−02.2	3460031	21:19:07.2	+46:18:48.0	18	799.38	0.038	N	Y	N
089.8−05.1	3460032	21:32:31.0	+44:35:47.7	18	799.38	0.035	N	N	N
095.2+00.7	3460033	21:31:50.2	+52:33:52.0	18	799.38	0.042	N	Y	N
100.6−05.4	3460034	22:23:55.7	+50:58:00.0	18	799.38	0.035	N	N	N
111.8−02.8	3460035	23:26:14.9	+58:10:53.0	8	355.28	0.038	N	N	N
118.0−08.6	3460096	00:18:42.2	+53:52:20.0	16	710.56	0.035	N	Y	N
123.6+34.5	3460115	12:33:06.8	+82:33:50.1	18	799.38	0.065	N	Y	N
146.7+07.6	3460097	04:25:50.8	+60:07:12.7	18	799.38	0.035	N	N	N
159.0−15.1	3460098	03:47:33.0	+35:02:48.9	9	399.69	0.045	N	Y	N
166.1+10.4	3460116	05:56:23.9	+46:06:17.2	18	799.38	0.050	N	N	N
190.3−17.7	3460099	05:05:34.3	+10:42:23.8	9	399.69	0.045	N	N	N
194.2+02.5	3460117	06:25:57.2	+17:47:27.3	18	799.38	0.045	N	N	N
211.2−03.5	3460036	06:35:45.1	+00:05:38.0	16	710.56	0.035	N	N	N
221.3−12.3	3460118	06:21:42.7	−12:59:14.0	15	666.15	0.046	N	N	N
226.7+05.6	3460100	07:37:18.9	−09:38:50.0	17	754.97	0.040	N	Y	N
232.8−04.7	3460037	07:11:16.7	−19:51:04.0	14	621.74	0.035	N	N	N
235.3−03.9	3460038	07:19:21.5	−21:43:54.3	17	754.97	0.035	N	N	N
258.1−00.3	3460039	08:28:28.0	−39:23:40.0	17	754.97	0.042	N	N	N
264.4−12.7	3460040	07:47:20.0	−51:15:03.4	17	754.97	0.035	N	Y	N
268.4+02.4	3460042	09:16:09.6	−45:28:42.8	18	799.38	0.038	N	N	N
278.6−06.7	3460101	09:19:27.5	−59:12:00.3	18	799.38	0.040	N	Y	N
283.8+02.2	3460044	10:31:33.4	−55:20:50.5	16	710.56	0.045	N	Y	N
285.4−05.3	3460120	10:09:20.8	−62:36:48.5	18	799.38	0.050	N	Y	N
285.6−02.7	3460045	10:23:09.1	−60:32:42.3	17	754.97	0.042	N	Y	N
285.7−14.9	3460121	09:07:06.3	−69:56:30.6	17	754.97	0.060	N	N	N
291.6−04.8	3460046	11:00:20.0	−65:14:57.8	18	799.38	0.038	N	Y	N
292.8+01.1	3460102	11:28:47.4	−60:06:37.3	18	799.38	0.042	N	Y	N
294.9−04.3	3460047	11:31:45.4	−65:58:13.7	18	799.38	0.035	N	N	N
296.3−03.0	3460048	11:48:38.2	−65:08:37.3	15	666.15	0.035	N	N	N
304.5−04.8	3460051	13:08:47.3	−67:38:37.6	16	710.56	0.038	N	Y	N
305.1+01.4	3460122	13:09:36.4	−61:19:35.6	18	82.44	0.050	Y	N	N
307.2−09.0	3460052	13:45:22.4	−71:28:55.7	18	799.38	0.035	N	Y	Y
307.5−04.9	3460053	13:39:35.1	−67:22:51.7	18	799.38	0.040	N	Y	Y
312.6−01.8	3460123	14:18:43.3	−63:07:10.1	16	710.56	0.056	N	Y	N
315.1−13.0	3460054	15:37:11.2	−71:54:52.9	18	82.44	0.042	Y	N	N
320.1−09.6	3460103	15:56:01.7	−66:09:09.2	16	710.56	0.045	N	Y	N
320.9+02.0	3460056	15:05:59.2	−55:59:16.5	9	399.69	0.042	N	Y	N
322.5−05.2	3460060	15:47:41.2	−61:13:05.6	16	710.56	0.092	N	Y	Y
323.9+02.4	3460061	15:22:19.4	−54:08:13.1	9	399.69	0.045	N	Y	N
324.8−01.1	3460062	15:41:58.8	−56:36:25.6	8	355.28	0.036	N	Y	Y
325.8−12.8	3460063	16:54:35.2	−64:14:28.4	16	710.56	0.035	N	Y	N
326.0−06.5	3460064	16:15:42.3	−59:54:01.0	14	621.74	0.035	N	Y	Y
327.1−01.8	3460065	15:58:08.1	−55:41:50.3	9	399.69	0.035	N	Y	Y

**Table 1**  
(Continued)

PN G	Obs. ID	R.A. (J2000) (hh:mm:ss)	Decl. (J2000) (dd:mm:ss)	$N^a$	$t_{\text{int}}^b$ (s)	FWHM ( $\mu\text{m}$ )	Flags		
							$\mathcal{S}^c$	$\mathcal{D}^d$	$\mathcal{C}^e$
327.8−01.6	3460066	16:00:59.1	−55:05:39.7	8	355.28	0.038	N	Y	N
331.1−05.7	3460067	16:37:42.7	−55:42:26.5	17	754.97	0.035	N	Y	N
331.3+16.8	3460068	15:12:50.8	−38:07:32.0	16	710.56	0.048	N	Y	N
336.3−05.6	3460069	16:59:36.1	−51:42:08.4	18	799.38	0.035	N	Y	N
342.1+27.5	3460104	15:22:19.3	−23:37:32.0	18	799.38	0.050	N	Y	N
349.8+04.4	3460105	17:01:06.2	−34:49:38.0	9	399.69	0.040	N	Y	N
350.9+04.4	3460070	17:04:36.2	−33:59:18.0	18	799.38	0.038	N	Y	N
356.1+02.7	3460076	17:25:19.2	−30:40:41.0	18	799.38	0.035	N	Y	Y
357.6+02.6	3460082	17:29:42.7	−29:32:50.0	15	666.15	0.035	N	Y	N

**Notes.**<sup>a</sup> The number of exposures.<sup>b</sup> The net integration time.<sup>c</sup> The saturation flag, Y: the spectrum is (partly) saturated in long-exposure spectral images; N: the spectrum is not saturated.<sup>d</sup> The decomposition flag, Y: the spectral decomposition procedure was used; N: the spectrum was not affected by any neighbor object.<sup>e</sup> The contamination flag, Y: the decontamination procedure was used; N: the spectrum was not affected by any overlapping object.**Table 2**  
Miscellaneous Information of Objects

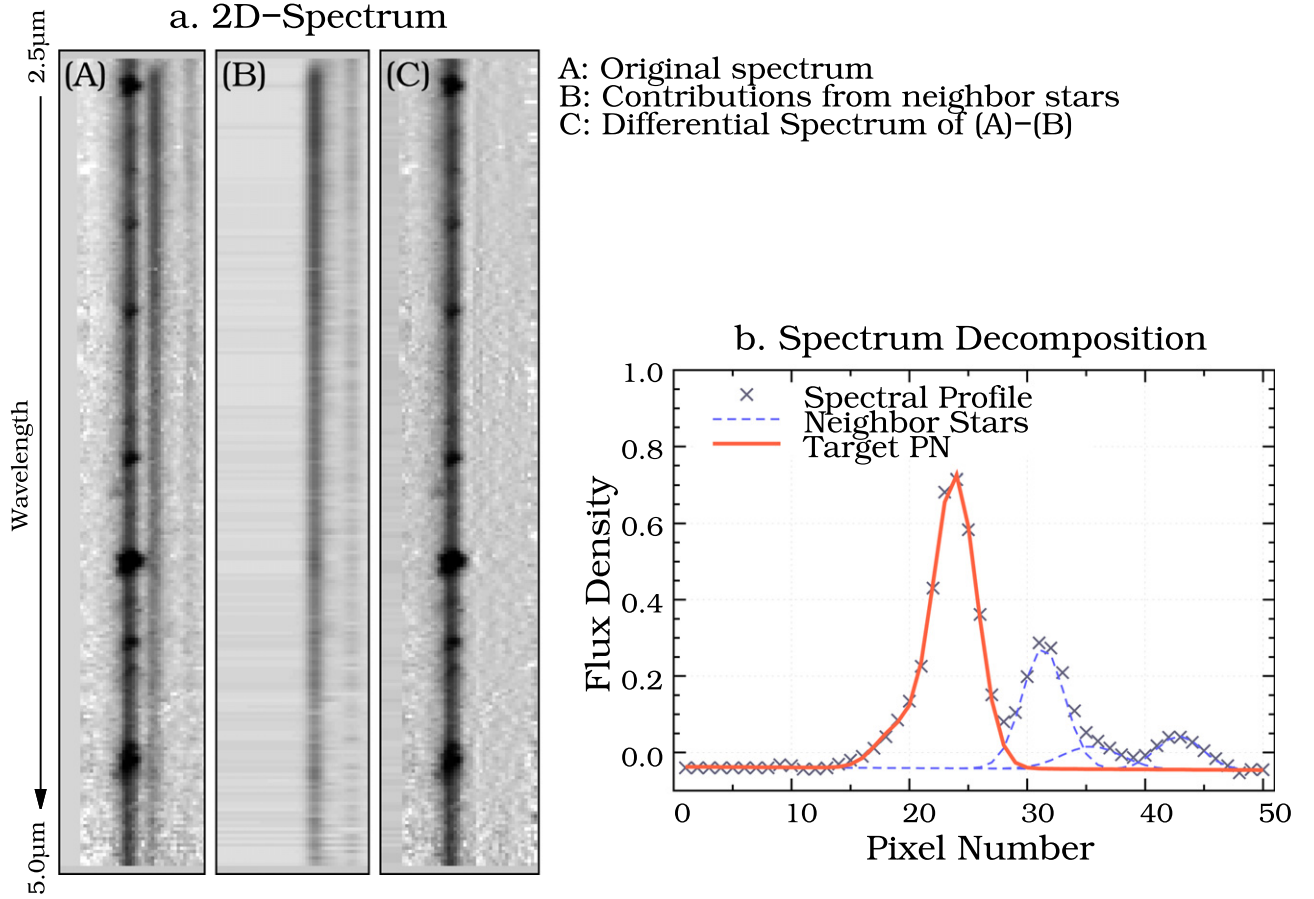
PN G	$V$ (a) (mag)	$K_s$ (b) (mag)	$A_V$ (H $\beta$ ) (mag)	$A_V$ (fit.) (mag)	$T_{\text{eff}}$ ( $10^3$ K)	References $T_{\text{eff}}$
000.3+12.2	13.9	11.4	$0.90^{+0.03}_{-0.03}$	$0.98^{+0.15}_{-0.05}$	47	Ka76, Ph03
002.0−13.4	14.1	11.1	$0.31^{+0.02}_{-0.01}$	$0.38^{+0.02}_{-0.04}$	60	PM89, PM91, Ph03
003.1+02.9	>17.0	11.6	$3.31^{+0.03}_{-0.03}$	$3.4^{+0.6}_{-0.1}$	87	PM89, Ph03
011.0+05.8	...	11.8	$1.81^{+0.04}_{-0.04}$	$1.79^{+0.06}_{-0.06}$	93	PM91, Ph03
027.6−09.6	15.2	11.9	$1.04^{+0.06}_{-0.05}$	$1.0^{+0.9}_{-0.2}$	58	PM91, Ka76, KJ91, Ph03
037.8−06.3	...	9.47	$1.4^{+0.1}_{-0.1}$	$1.41^{+0.08}_{-0.02}$	74	Ph03
038.2+12.0	12.5	11.1	$0.82^{+0.05}_{-0.05}$	$0.93^{+0.07}_{-0.07}$	30	Ka78, Ph03
043.1+03.8	14.9	12.1	$1.9^{+0.1}_{-0.1}$	$2.05^{+0.46}_{-0.05}$	28	Ph03
046.4−04.1	15.2	11.1	$1.42^{+0.02}_{-0.02}$	$1.39^{+0.01}_{-0.01}$	68	Ka76, KJ91, Ph03
051.4+09.6	13.3	10.4	$0.88^{+0.05}_{-0.04}$	$0.98^{+0.11}_{-0.08}$	37	Ka76, Ka78, Ph03

**References.** (a) Acker et al. (1992), (b) Skrutskie et al. (2006), HF83: Harrington & Feibelman (1983), Ka76: Kaler (1976), Ka78: Kaler (1978), KJ91: Kaler & Jacoby (1991), Lu01: Lumsden et al. (2001), Me88: Mendez et al. (1988), Ph03: Phillips (2003), PM89: Preite-Martinez et al. (1989), PM91: Preite-Martinez et al. (1991).

region. The horizontal and vertical axes correspond to the position and wavelength. The spectrum of the target PN runs through the center of the panel. The spectra of neighboring objects are located to the right side of the PN spectrum. Figure 2(b) shows the spatial profiles of the spectra (Panel (A)). The spectra were decomposed by fitting with a combination of Gaussian functions. The red solid line shows the PN profile, and the blue dashed lines show those of the neighboring objects. The fitting was performed for every spectral element. The instrumental spatial profile of a spectrum was not a Gaussian, rather a combination of Gaussian profiles provided the best fit to the profile. Panel (B) of Figure 2(a) shows the estimated contribution from the neighboring objects. By subtracting Panel (B) from Panel (A), the two-dimensional spectrum of the target PN was extracted (Panel (C)). Finally, the one-dimensional spectrum of the target PN was obtained from Panel (C) with a wide aperture, as mentioned above. The uncertainty in the contamination-corrected spectrum was estimated by taking into account errors in the spectral decomposition procedure. The targets which are in crowded regions are denoted in the ninth column of Table 1.

**3.2.1. Subtraction of Contamination**

For some observations, the targets were found to be aligned with neighboring objects, along the direction of dispersion. In such cases, the spectra of these objects fully overlapped with each other. This made retrieval of the target profile spectrum, by profile fitting alone, impossible (see the upper panels in Figure 3). To remove the contamination from the overlapping objects, we developed a method to estimate the amount of the contamination (hereafter the decontamination procedure). The method basically follows the procedure used in Sakon et al. (2008). At first, a one-dimensional spectrum was extracted with a wide aperture. We assumed that overlapping objects were stars and their spectra did not show any particular spectral features in the spectral range in question. Their 2.5–5.0  $\mu\text{m}$  spectra ( $\tilde{f}_\nu$ ) were estimated from the 2MASS, AKARI, and WISE photometric data by fitting with a blackbody function. The calibration of the overlapping objects required correction since the overlapping object was shifted along the direction of dispersion. The amount of contamination  $f_\nu^c(\lambda)$  was converted from  $\tilde{f}_\nu$  using the distance between the target and the overlapping objects in the direction of the dispersion,  $d$ , in



**Figure 2.** Schematic view of the spectral decomposition procedure. (Left) Panels (A), (B), and (C) represent an original two-dimensional spectrum, estimated contributions from the neighboring stars, and an isolated target spectrum. (Right): spatial profiles of the two-dimensional spectrum (Panel (A) of Figure 2(a)). The profiles are decomposed by fitting (see text). The contributions from neighboring stars are shown by the blue dashed lines. The component of the target is shown by the red solid line.

units of pixel and the spectral response function of the IRC,  $R(\lambda)$ , in units of  $\text{ADU mJy}^{-1}$ . We defined  $\delta$  as the pixel scale of the spectrum in units of  $\mu\text{m pixel}^{-1}$ . The amount of the contamination was estimated by

$$f_{\nu}^c(\lambda) = \frac{\tilde{f}_{\nu}(\lambda - \delta d)R(\lambda - \delta d)}{R(\lambda)}. \quad (1)$$

We defined  $f_{\nu}^{\text{obs}}$  as the obtained spectrum, including the contamination. The decontaminated spectrum was defined by  $f_{\nu}^{\text{obj}} = f_{\nu}^{\text{obs}} - f_{\nu}^c$ . A schematic view of the method is given in Figure 3. The upper left panel shows the location of the target (PN G123.6+34.5) and overlapping object. The direction of dispersion is horizontal. The spectrum with the contamination is shown in the upper right panel. The contaminated one-dimensional spectrum is shown by the gray dotted line in the bottom panel. The estimated contamination and decontaminated spectra are shown by the red dashed and blue solid lines. We assume that the spectrum of the overlapping star does not have any significant features and that the decontamination procedure does not affect the estimate of emission feature intensities (although continuum flux could be somewhat affected). The targets with overlapping objects are flagged as contaminated objects in the tenth column of Table 1.

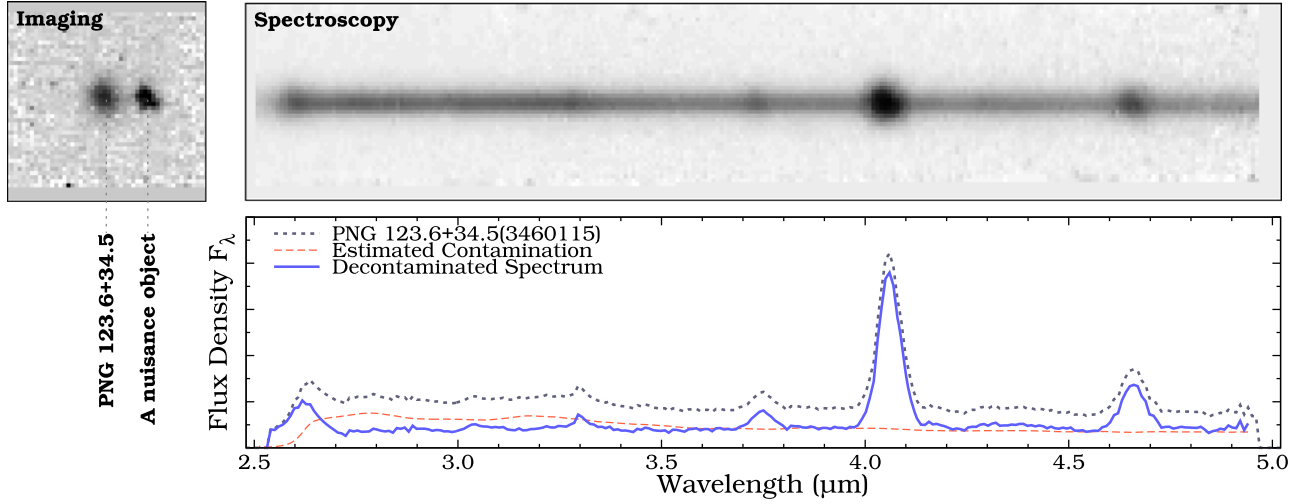
## 4. RESULTS

### 4.1. Near-infrared Spectroscopy of Planetary Nebulae Spectral Atlas

Details of the observations are summarized in Table 1, which includes the PNG ID, the observation ID, the location, the number of exposures, the total integration time, the spectral PSF size, and the flags to denote data quality. The PNSPC catalog contains extracted one-dimensional spectra shown in Figure 7. Most of the spectra cover 2.5–5.0  $\mu\text{m}$ . Some spectra were affected by scattered light inside the instrument or column-pull-down effects of the detector, which were not correctable at present. The part of the spectrum severely affected by such issues has been excluded from the catalog.

Two representative spectra in the PNSPC catalog are shown in Figure 4. The top panel is the spectrum of PNG 064.7+05.0, and the bottom panel is that of PNG 074.5+02.1. The spectra showed several emission features: H I and He II RLs, the PAH feature at 3.3  $\mu\text{m}$ , and the aliphatic feature complex in 3.4–3.5  $\mu\text{m}$ . Some spectra also showed He II RLs and [Mg IV] at 4.49  $\mu\text{m}$  and [Ar VI] at 4.53  $\mu\text{m}$  fine-structure lines, indicating that the excitation of the object was high. Some objects showed molecular hydrogen lines, but they were weak and not frequently detected. No prominent absorption features were detected in any of the objects.

The top panel of Figure 5 shows a histogram of the uncertainty in continuum emission at 3.0  $\mu\text{m}$ . The gray dashed



**Figure 3.** Schematic view of the decontamination procedure. The upper left panel shows the locations of a PN and an overlapping star, while the upper right panel shows the obtained spectrum with contamination from the overlapping star. The contaminated spectrum is shown by the gray dotted line in the bottom panel. The red dashed line shows the estimated contamination from the overlapping star. The decontaminated spectrum of the PN is shown by the blue solid line.

line indicates the median of the uncertainty, which is about  $0.7 \times 10^{-15} \text{ W m}^{-2} \mu\text{m}^{-1}$  or 2 mJy. When the object is brighter than  $2\text{--}5 \times 10^{-13} \text{ W m}^{-2} \mu\text{m}^{-1}$ , the pixels become saturated in the long exposure image. For moderately bright objects, pixels around emission lines were saturated. The spectra of extremely bright objects were completely saturated in the long exposure images. The S/N of the spectrum depends on the fraction of saturated pixels. PNG 064.7+05.0 and 305.1+01.4 are extremely bright, with their spectrum completely saturated in the long exposure images. Figure 5 (top) indicates that their uncertainties at  $3.0 \mu\text{m}$  are about 50 times worse than the others. The fractional uncertainty at  $3.0 \mu\text{m}$  is shown in the bottom panel of Figure 5. The median of the fractional uncertainty is about 6.4%, indicated by the vertical dashed line. This means that, for half of the PNSPC samples, the continuum emission at  $3.0 \mu\text{m}$  was detected with a S/N larger than about 15.

The FWHM of the spectral PSF of the IRC, equivalent to the FWHM of an unresolved emission line, is about  $0.035 \mu\text{m}$ . Thus, the spectral resolution is about  $\lambda/\Delta\lambda \sim 100$ . When the object is extended for the IRC, the spectral resolution is degraded. The FWHM of the spectral PSF, estimated from spectral fitting (see Section 4.3), is listed in the seventh column of Table 1. PN G322.5–05.2 is the most extended object in the PNSPC catalog and its spectral resolution was estimated to be about  $\lambda/\Delta\lambda \sim 40$  at  $4 \mu\text{m}$ .

#### 4.2. Extinction

Foreground extinction toward the objects was estimated based on an intensity ratio of  $\text{H}\beta$  to Brackett- $\alpha$  ( $\text{Br}\alpha$ ). The intensities of  $\text{H}\beta$  were obtained from literature (Acker et al. 1992, and references therein). We measured the intensities of  $\text{Br}\alpha$  from the *AKARI* spectra by fitting with a Gaussian function. We defined  $I_{\text{H}\beta}$  and  $I_{\text{Br}\alpha}$  as the intensities of  $\text{H}\beta$  and  $\text{Br}\alpha$ . The observed intensity ratio follows the relation:

$$\frac{I_{\text{H}\beta}^{\text{obs}}}{I_{\text{Br}\alpha}^{\text{obs}}} = \frac{I_{\text{H}\beta}^0 e^{-\tau_{\text{H}\beta}}}{I_{\text{Br}\alpha}^0 e^{-\tau_{\text{Br}\alpha}}}, \quad (2)$$

where  $I_X^{\text{obs}}$  and  $I_X^0$  are the observed and intrinsic intensities of the line  $X$ , and  $\tau_X$  is the extinction of the line  $X$ . We assumed

that the nebula was totally opaque for Lyman photons (Case B in Baker & Menzel 1938). Then, the intrinsic intensity ratio ( $I_{\text{H}\beta}^0/I_{\text{Br}\alpha}^0$ ) was assumed to be  $\sim 12.853$ , given by the Case B line ratio calculated for the electron density of  $10^4 \text{ cm}^{-3}$  and the electron temperature of  $10^4 \text{ K}$  from Storey & Hummer (1995). We adopted the extinction curve given by Mathis (1990) and the extinction at the  $V$ -band was derived as

$$A_V(\text{H}\beta) = 2.39 - 0.95 \ln \left( \frac{I_{\text{H}\beta}^{\text{obs}}}{I_{\text{Br}\alpha}^{\text{obs}}} \right). \quad (3)$$

The coefficients in Equation (3) depends on the assumed electron temperature and density. When the assumed electron temperature increases to  $2 \times 10^4 \text{ K}$ ,  $A_V$  will increase by about 0.3. The effect of the electron density on  $A_V$  is much smaller than 0.3 for the range from  $10^2$  to  $10^6 \text{ cm}^{-3}$ . The systematic uncertainty in  $A_V$ , chiefly attributable to the electron temperature uncertainty, was estimated to be less than 0.3 mag. The estimated  $A_V(\text{H}\beta)$  is listed in the fourth column of Table 2.

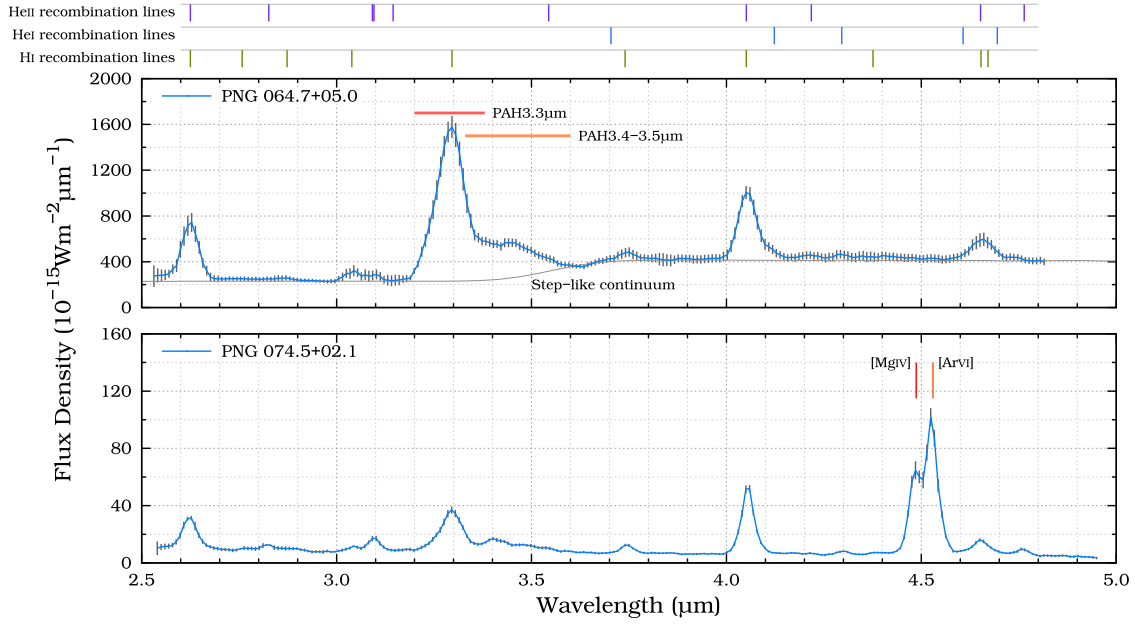
#### 4.3. Spectral Fitting

The intensities of the spectral features were measured by spectral fitting. The fitted function was assumed to be a linear combination of the spectral features as

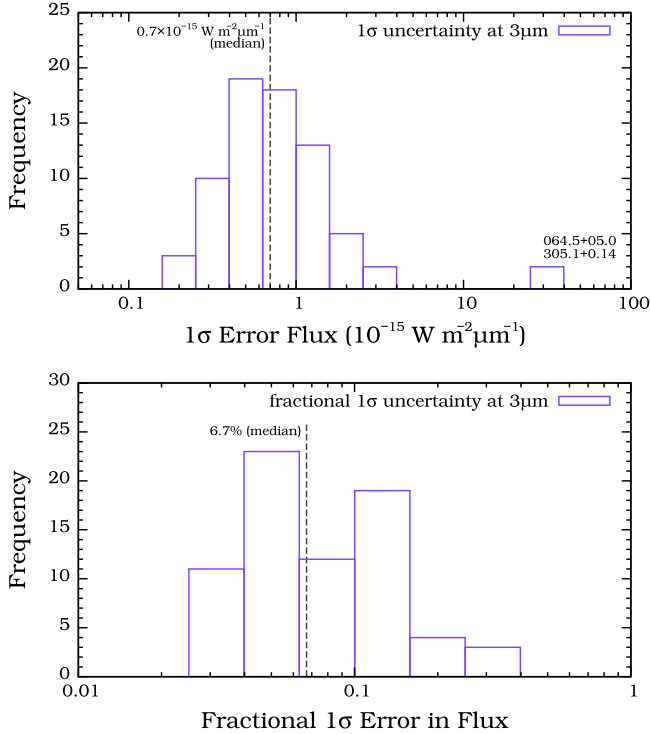
$$F_\lambda(\lambda) = [f_\lambda^{\text{cont}}(\lambda) + f_\lambda^{\text{line}}(\lambda) + f_\lambda^{\text{dust}}(\lambda)]e^{-\tau(\lambda)}, \quad (4)$$

where  $f_\lambda^{\text{cont}}$ ,  $f_\lambda^{\text{line}}$ , and  $f_\lambda^{\text{dust}}$  were the components of continuum, emission lines, and dust features, and  $\tau(\lambda)$  was the optical depth at  $\lambda$ . The functional shape of  $\tau(\lambda)$  was given by a cubic spline function that interpolated the extinction curve of Mathis (1990) at every data point of the IRC spectrum. The amount of the extinction at the  $V$ -band,  $A_V = \tau(V)$ , was included as a fitting parameter.

The continuum emission,  $f_\lambda^{\text{cont}}$ , was given by a polynomial function,  $f_\lambda^{\text{cont}} = \sum_{l=0}^N c_l \lambda^l$ , where  $c_l$ 's were the fitting parameters. The order  $N$  was defined for each object, usually  $N = 2$  or  $3$ . This was sufficient to represent the contributions from free-free and hot thermal dust emission. Some PNe



**Figure 4.** Representative spectra in the PNSPC catalog. The top and bottom panels show the spectra of PNG 064.7+05.0 and PNG 074.5+02.1 by the blue solid lines with errors. Series of recombination lines (H I, He I, and He II) are indicated by the vertical lines. The both spectra show the PAH feature in  $3.3 \mu\text{m}$  and the aliphatic feature complex in  $3.4\text{--}3.5 \mu\text{m}$ . The spectrum of PNG 074.5+02.1 shows fine-structure lines of [Mg IV] at  $4.487 \mu\text{m}$  and [Ar VI] at  $4.530 \mu\text{m}$ . Step-like continuum emission is seen in the spectra of PNG 064.7+05.0 (see Section 4.3).



**Figure 5.** Histograms of the uncertainty in continuum emission at  $3.0 \mu\text{m}$ . The top panel shows the histogram of absolute  $1\sigma$  errors. The gray dashed line indicates the median of the uncertainty ( $0.7 \times 10^{-15} \text{ W m}^{-2} \mu\text{m}^{-1}$ ). The data around  $30 \times 10^{-15} \text{ W m}^{-2} \mu\text{m}^{-1}$  are very bright sources of PNG 064.7+0.5 and 305.1+01.4, for which the spectra were entirely saturated in the long exposure images. The bottom panel shows the histogram of fractional  $1\sigma$  errors. The median value ( $\sim 6.7\%$ ) is shown by the gray dashed line.

showed a sudden increase in the continuum emission around  $\lambda \sim 3.6 \mu\text{m}$ , for instance, PNG 064.7+05.0 in Figure 4. The existence of this emission was mentioned in Boulanger et al. (2011), but the carrier of this emission has not been identified.

It was difficult to fit this profile with such a low-order polynomial function. Thus, for those objects with such an increase, we added a step-like function in the components of the continuum emission as

$$f_{\lambda}^{\text{cont}}(\lambda) = \sum_{l=0}^N c_l \lambda^l + c \left[ \frac{1}{2} + \frac{1}{2} \text{erf} \left( \frac{\lambda - 3.65 \mu\text{m}}{\Delta_{\text{step}}} \right) \right], \quad (5)$$

where  $c$  was the fitting parameter,  $\text{erf}(x)$  was the error function, and  $\Delta_{\text{step}}$  was the width of the step function, which was tentatively fixed at  $0.135 \mu\text{m}$ . This component strongly appears in PNG 037.8–06.3, 064.7+05.0, 089.8–05.1, and 268.4+02.4.

The emission line profile was approximated by a Gaussian function. The line emission component,  $f_{\lambda}^{\text{line}}$ , was approximated by a combination of Gaussian functions:

$$f_{\lambda}^{\text{line}}(\lambda) = \sum_m \frac{c_m}{\sqrt{2\pi}\sigma} \exp \left[ -\frac{(\lambda - \lambda_{m0})^2}{2\sigma^2} \right], \quad (6)$$

where  $\lambda_{m0}$  was the central wavelengths of the  $m$ th emission line and  $c_m$ 's were the parameters to indicate the strength of the lines. The central wavelengths were fixed in the fitting. The line width of the Br $\alpha$  line was measured by fitting with a Gaussian function prior to the fitting and the line widths of other emission lines were assumed to be the same as Br $\alpha$ . The lines included in the fitting are listed in Table 3. The assignment and central wavelength are shown in the first and second column. In the fitting, we did not include the lines which were apparently not detected. Some of the H I and He II RLs appear at the same wavelength and it was impossible to isolate them. Thus, the relative intensities of the H I and He II RLs were fixed, assuming the Case B condition. The relative intensity was adopted from Storey & Hummer (1995) for the electron temperature and the electron density of  $10^4 \text{ K}$  and  $10^4 \text{ cm}^{-3}$ . The variations in the relative intensity with the electron



**Table 3**  
List of Line Features

Name	$\lambda_0$ ( $\mu\text{m}$ )
Brakett- $\alpha^a$	4.051
Brakett- $\beta^a$	2.625
Pfund- $\beta^a$	4.653
Pfund- $\gamma^a$	3.740
Pfund- $\delta^a$	3.296
Pfund- $\epsilon^a$	3.039
Pfund- $\zeta^a$	2.873
Pfund- $\eta^a$	2.758
Humpleys- $\epsilon^a$	4.671
Humpleys- $\zeta^a$	4.376
He I(3P <sup>o</sup> -3D)	3.704
He I(1P <sup>o</sup> -1D)	4.123
He I(3S-3P <sup>o</sup> )	4.296
He I(1P <sup>o</sup> -1S)	4.607
He I(3P <sup>o</sup> -3S)	4.695
He II(7-6) <sup>a</sup>	3.092
He II(9-7) <sup>a</sup>	2.826
He II(8-7) <sup>a</sup>	4.764
He II(12-8) <sup>a</sup>	2.625
He II(11-8) <sup>a</sup>	3.096
He II(10-8) <sup>a</sup>	4.051
He II(14-9) <sup>a</sup>	3.145
He II(13-9) <sup>a</sup>	3.544
He II(12-9) <sup>a</sup>	4.218
He II(14-10) <sup>a</sup>	4.652
H <sub>2</sub> (1-0)	2.802
H <sub>2</sub> (1-0)	3.003
H <sub>2</sub> (1-0)	3.234
H <sub>2</sub> (0-0)	3.836
H <sub>2</sub> (0-0)	4.181
[Mg IV]	4.487
[Ar VI]	4.530

**Note.**

<sup>a</sup> Relative intensities of these lines are fixed assuming the Case B condition.

temperature and density were almost negligible compared to the uncertainty in the observed spectra.

The intrinsic spectral profiles of the PAH emission and the aliphatic feature complex were assumed to be given by a combination of Lorentzian functions which need to be convoluted with the resolution given by the instrument. Thus, their spectral profiles were approximated by the Lorentzian function convoluted with the spectral PSF of the IRC, which was approximated by a Gaussian. The component from dust emission,  $f_{\lambda}^{\text{dust}}$ , was defined by

$$f_{\lambda}^{\text{dust}}(\lambda) = \sum_n \frac{c_n}{\pi \sqrt{2\pi} \sigma} \int d\lambda' e^{-\frac{(\lambda-\lambda')^2}{2\sigma^2}} \frac{\Delta_n}{\Delta_n^2 + (\lambda' - \lambda_{n0})^2}, \quad (7)$$

where  $\Delta_n$  and  $\lambda_{n0}$  were the width and peak wavelength of the  $n$ th dust emission and  $c_n$ 's were the fitting parameters. The adopted values of  $\Delta_n$ 's and  $\lambda_{n0}$ 's were listed in Table 4. The PAH feature at 3.3  $\mu\text{m}$  was mainly composed of the 3.30  $\mu\text{m}$  feature, but the 3.25  $\mu\text{m}$  feature was added to account for a slight change in the spectral profile of the 3.3  $\mu\text{m}$  feature. For instance the 3.3  $\mu\text{m}$  feature of PNG 320.9+02.0 show an extended blue component, which cannot be well-fitted without the 3.25  $\mu\text{m}$  component. Mori et al. (2012) used a combination of four Lorentzian functions to describe the 3.4–3.5  $\mu\text{m}$  feature complex (C–H<sub>al</sub>) in spectra obtained with the IRC. We used a

**Table 4**  
List of Dust Features

Name	$\lambda_0$ ( $\mu\text{m}$ )	$\Delta$ ( $\mu\text{m}$ )
PAH 3.25 <sup>a</sup>	3.248	0.005
PAH 3.30 <sup>a</sup>	3.296	0.021
C–H <sub>al</sub> 3.41 <sup>b</sup>	3.410	0.030
C–H <sub>al</sub> 3.46 <sup>b</sup>	3.460	0.013
C–H <sub>al</sub> 3.51 <sup>b</sup>	3.512	0.013
C–H <sub>al</sub> 3.56 <sup>b</sup>	3.560	0.013

**Notes.**

<sup>a</sup> They belong to the PAH feature at 3.3  $\mu\text{m}$ .

<sup>b</sup> They belong to the aliphatic feature complex in 3.4–3.5  $\mu\text{m}$ .

combination of the four components to describe the C–H<sub>al</sub> feature. The central wavelengths and widths were adopted from Mori et al. (2012).

Define  $F_{\lambda}^{\text{obj}}(\lambda)$  and  $\sigma_{\lambda}(\lambda)$  as the observed flux and its uncertainty at  $\lambda$ . We assumed that the residual  $\Delta_{\lambda}(\lambda) = F_{\lambda}^{\text{obj}}(\lambda) - F_{\lambda}(\lambda)$  followed a normal distribution with the variance of  $\sigma_{\lambda}^2(\lambda)$ . The likelihood function was defined as

$$\mathcal{L} \propto \prod_i \exp \left[ -\frac{1}{2} \left( \frac{\Delta_{\lambda}(\lambda_i)}{\sigma_{\lambda}(\lambda_i)} \right)^2 \right], \quad (8)$$

where  $\lambda_i$  was the wavelength of the  $i$ th spectral element. The fitting parameters could be defined by maximizing Equation (8). However, it was difficult to accurately derive the amount of the extinction,  $A_V$ , only from the near-infrared spectrum (see, Appendix). To make the fitting stable, we used the extinction estimated in Section 4.2 to provide a prior distribution of the extinction

$$P(A_V) \propto \exp \left[ -\frac{1}{2} \left( \frac{A_V - A_V(\text{H}\beta)}{\sigma''} \right)^2 \right], \quad (9)$$

where  $A_V(\text{H}\beta)$  was the extinction estimated from the H $\beta$  to Br $\alpha$  intensity ratio and  $\sigma''$  was tentatively fixed at 0.3 mag, estimated from the systematic error. The fitting parameters were derived by maximizing the product of Equations (8) and (9) with the constraints that the parameters of  $A_V$ ,  $c$ ,  $c_m$  and  $c_n$  should be non-negative. The uncertainties of the fitting parameters were estimated based on a Monte Carlo simulation: The spectrum of the target,  $\tilde{F}(\lambda)$  was simulated assuming a normal distribution of flux with a mean of  $F_{\lambda}^{\text{obj}}(\lambda)$  and a variance of  $\sigma_{\lambda}(\lambda)$ . The fitting parameters were derived for  $\tilde{F}(\lambda)$ . This procedure was repeated 1000 times and the probability distributions of the fitting parameters were obtained. The confidence interval of the fitting parameters was defined to account for 68% of the trials.

The extinction obtained by the fitting ( $A_V(\text{fit.})$ ) is listed in the fifth column of Table 2. Table 5 lists the extinction-corrected intensities of Br $\alpha$  at 4.05  $\mu\text{m}$ , He I at 4.30  $\mu\text{m}$ , He II at 3.09  $\mu\text{m}$ , the PAH feature at 3.3  $\mu\text{m}$ , the aliphatic feature complex in 3.4–3.5  $\mu\text{m}$ , [Mg IV] at 4.49  $\mu\text{m}$ , and [Ar VI] at 4.53  $\mu\text{m}$ . These emission features were strong and widely detected among the PNSPC samples. Other line features were weak or not detected in the majority of the targets and thus were not included in the table. The equivalent widths of those are listed in Table 6. Note that [Mg IV] and [Ar VI] are so close

**Table 5**  
Extinction Corrected Line Intensity

PN G	Bracket- $\alpha$ ( $10^{-15} \text{ W m}^{-2}$ )	He II (3.09 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )	He I (4.30 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )	PAH (3.3 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )	C-H <sub>al</sub> (3.4–3.5 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )	[Mg IV] (4.49 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )	[Ar VI] (4.53 $\mu\text{m}$ ) ( $10^{-15} \text{ W m}^{-2}$ )
000.3+12.2	2.74 $^{+0.05}_{-0.04}$	...	0.14 $^{+0.03}_{-0.03}$	<0.11	...	...	...
002.0–13.4	2.05 $^{+0.04}_{-0.03}$	...	0.31 $^{+0.03}_{-0.02}$	0.25 $^{+0.05}_{-0.05}$	<0.10	...	...
003.1+02.9	3.16 $^{+0.04}_{-0.04}$	<0.53	0.43 $^{+0.03}_{-0.04}$	0.94 $^{+0.25}_{-0.16}$	0.49 $^{+0.09}_{-0.08}$	0.35 $^{+0.05}_{-0.04}$	...
011.0+05.8	1.09 $^{+0.03}_{-0.04}$	0.13 $^{+0.03}_{-0.03}$	<0.13	...	...	1.11 $^{+0.05}_{-0.08}$	0.20 $^{+0.04}_{-0.05}$
027.6–09.6	1.12 $^{+0.04}_{-0.02}$	...	0.10 $^{+0.01}_{-0.01}$	...	...	...	...
037.8–06.3	4.61 $^{+0.24}_{-0.10}$	...	1.80 $^{+0.13}_{-0.09}$	6.64 $^{+0.40}_{-0.21}$	2.60 $^{+0.21}_{-0.14}$	...	...
038.2+12.0	2.18 $^{+0.04}_{-0.06}$	...	<0.13	0.45 $^{+0.04}_{-0.05}$	0.12 $^{+0.03}_{-0.02}$	...	...
043.1+03.8	0.45 $^{+0.01}_{-0.01}$	...	...	<0.06	<0.01	...	...
046.4–04.1	2.43 $^{+0.02}_{-0.01}$	0.11 $^{+0.02}_{-0.02}$	0.32 $^{+0.02}_{-0.01}$	0.40 $^{+0.04}_{-0.05}$	<0.16	...	...
051.4+09.6	3.25 $^{+0.06}_{-0.02}$	...	0.52 $^{+0.02}_{-0.02}$	0.97 $^{+0.05}_{-0.07}$	0.23 $^{+0.04}_{-0.05}$	...	...

**Table 6**  
Extinction Corrected Line Equivalent Width

PN G	Bracket- $\alpha$ ( $\text{\AA}$ )	He II (3.09 $\mu\text{m}$ ) ( $\text{\AA}$ )	He I (4.30 $\mu\text{m}$ ) ( $\text{\AA}$ )	PAH (3.3 $\mu\text{m}$ ) ( $\text{\AA}$ )	C-H <sub>al</sub> (3.4–3.5 $\mu\text{m}$ ) ( $\text{\AA}$ )	[Mg IV] (4.49 $\mu\text{m}$ ) ( $\text{\AA}$ )	[Ar VI] (4.53 $\mu\text{m}$ ) ( $\text{\AA}$ )
000.3+12.2	2825.6 $^{+47.3}_{-38.2}$	...	154.5 $^{+35.5}_{-32.7}$	<130.7	...	...	...
002.0–13.4	2938.4 $^{+54.0}_{-39.9}$	...	481.9 $^{+42.1}_{-27.1}$	435.3 $^{+86.8}_{-84.6}$	<181.4	...	...
003.1+02.9	2105.7 $^{+26.5}_{-25.5}$	<236.7	343.7 $^{+25.6}_{-30.2}$	665.7 $^{+176.0}_{-113.7}$	365.1 $^{+68.0}_{-59.5}$	324.3 $^{+51.4}_{-42.0}$	...
011.0+05.8	2005.1 $^{+61.4}_{-68.8}$	144.8 $^{+32.1}_{-34.2}$	<266.9	...	...	2605.4 $^{+127.9}_{-179.7}$	486.2 $^{+101.8}_{-111.4}$
027.6–09.6	3229.6 $^{+128.8}_{-44.8}$	...	328.8 $^{+35.8}_{-27.6}$	...	...	...	...
037.8–06.3	482.1 $^{+25.4}_{-10.0}$	...	182.2 $^{+13.6}_{-8.6}$	1342.1 $^{+80.2}_{-42.1}$	531.1 $^{+42.6}_{-29.1}$	...	...
038.2+12.0	2863.3 $^{+54.5}_{-84.4}$	...	<192.5	709.0 $^{+55.0}_{-75.8}$	204.3 $^{+47.2}_{-38.7}$	...	...
043.1+03.8	1316.2 $^{+42.9}_{-37.5}$	...	...	<230.0	...	...	...
046.4–04.1	3171.2 $^{+29.3}_{-5.0}$	99.6 $^{+17.8}_{-18.0}$	475.1 $^{+29.6}_{-9.3}$	595.7 $^{+65.3}_{-75.1}$	<250.8	...	...
051.4+09.6	1882.4 $^{+36.4}_{-13.5}$	...	315.8 $^{+15.0}_{-12.1}$	740.7 $^{+41.7}_{-56.1}$	180.1 $^{+32.7}_{-42.6}$	...	...

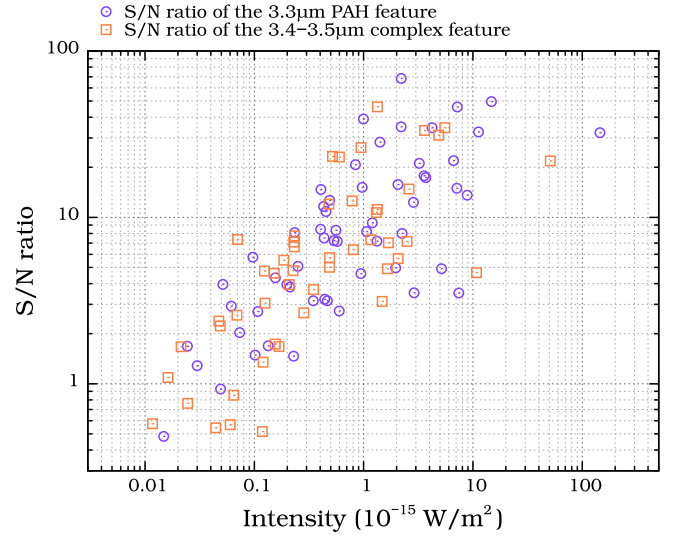
that their spectral profiles are in part overlapping. Their uncertainties are relatively large compared to other lines, because the effect of overlapping was taken into account. Figure 6 shows the S/N against the intensity of the 3.3  $\mu\text{m}$  PAH feature and the 3.4–3.5  $\mu\text{m}$  aliphatic feature complex by the blue circles and the red squares. Figure 6 indicates that the emission feature stronger than  $\sim 2 \times 10^{-16} \text{ W m}^{-2}$  was detected with a  $3\sigma$ -confidence level.

## 5. DISCUSSION

### 5.1. Biases in the PNSPC samples

#### 5.1.1. Galactic Coordinates

We compare the distribution of the PNSPC samples in the Galactic coordinates with that of the PNe in Acker et al. (1992), which are assumed to represent the distribution of all the PNe in the Milky Way. The histograms of the Galactic longitude and latitude are shown in the left and right panels of Figure 8. The histograms of the PNSPC samples are shown by the blue lines, while the histograms of PNe in Acker et al. (1992) are shown by the gray filled bars. The total number of the PNe in Acker et al. (1992) is 1143. The longitude histogram of the PNSPC samples looks broader than that of the PNe in Acker et al. (1992), while the latitude histogram of the PNSPC samples does not show a large deviation from that of the PNe in Acker et al. (1992). We investigate the differences by Kolmogorov–Smirnov test. The values of  $\chi^2_{\nu=2}$  are about 16 for the longitude and 4 for the latitude. The results suggest that the deviation of the longitude is highly significant and that the



**Figure 6.** The signal-to-noise ratio (S/N) is plotted against the intensity of the emission features, indicating a typical detection limit of emission features. The blue circles represent the S/N ratio of the 3.3  $\mu\text{m}$  PAH feature, while the red squares show that of the 3.4–3.5  $\mu\text{m}$  aliphatic feature complex.

PNSPC samples are biased toward PNe in the Galactic disk rather than those in the bulge. Figure 9 shows the position of the PNSPC samples projected onto the Galactic plane. The distances to the objects are taken from literature (Acker 1978, Pottasch 1983, p. 391; Amnuel et al. 1984; Maciel 1984; Gathier et al. 1986a, 1986b; Sabbadin 1986; Acker et al. 1992;

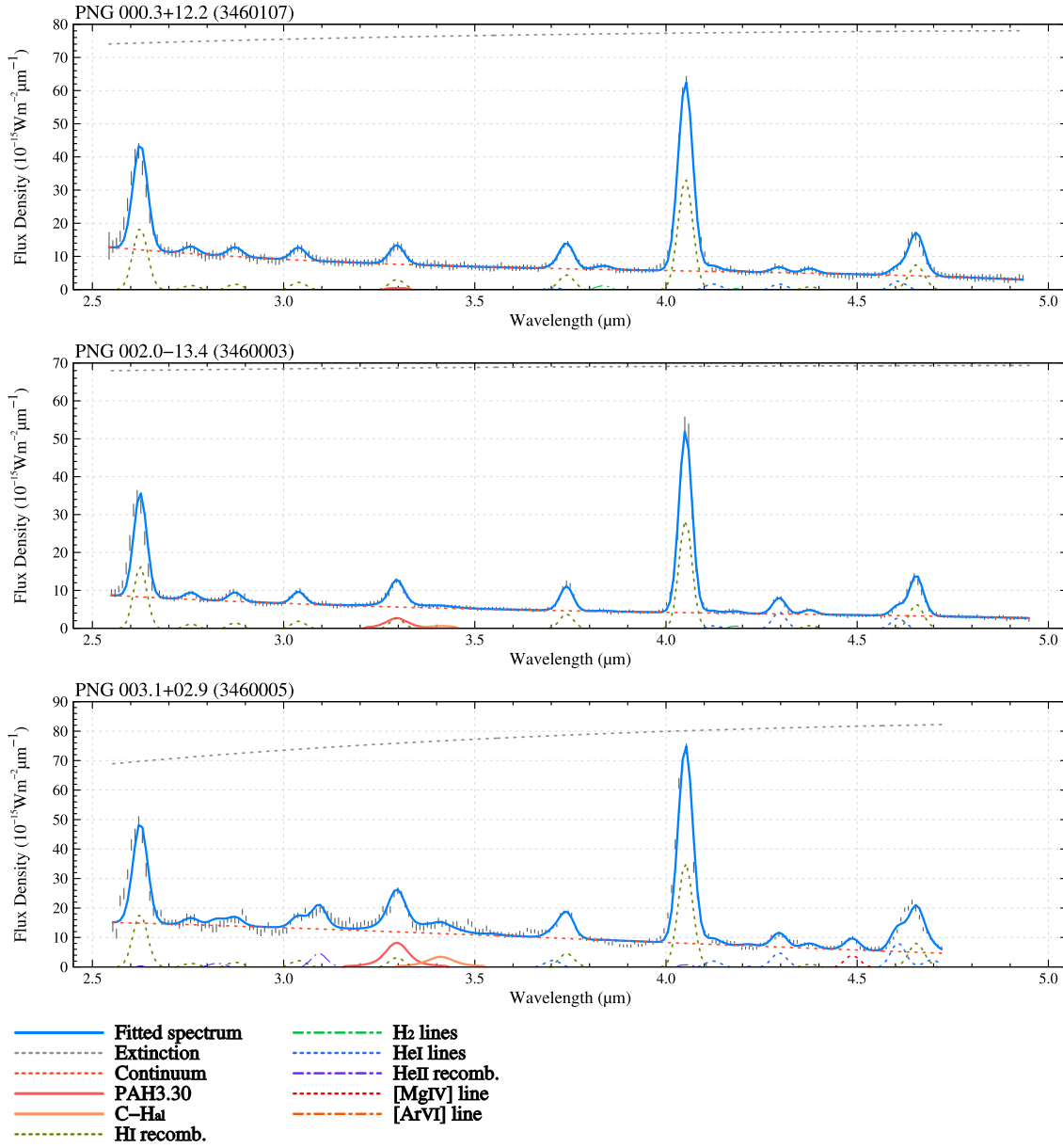


Figure 7. The AKARI/IRC 2.5–5.0  $\mu\text{m}$  spectra. Explanations of the lines are shown in the bottom of the figures.

Cahn et al. 1992, and references therein). Figure 9 shows that most of the PNSPC samples are within 5 kpc of the Sun, even when accounting for uncertainties in the distance estimate. This also supports that the PNSPC samples mainly consist of PNe in the Galactic disk.

### 5.1.2. Effective Temperature and Age

The central star of a PN becomes hot as it evolves. The effective temperature of the PN can be used as an indicator of the age of the PN. The timescale of the evolution depends on the mass of the central star (e.g., Schönberner 1983; Blöcker 1995). Although the individual age of the PN is difficult to accurately derive from the effective temperature, the distribution of the effective temperature reflects the bias in the age of the PNSPC samples. The effective temperatures were collected from literature. The temperature and its references are listed in the sixth and seventh columns of Table 2. The effective

temperature is estimated with several different methods. The effective temperatures estimated based on a model atmosphere should be the most reliable. The temperatures estimated in Mendez et al. (1988) and Harrington & Feibelman (1983) are preferentially used, if available. The typical error for temperatures measured by this method is about 10% (Mendez et al. 1988). Otherwise, the effective temperatures estimated with the Zanstra or the Energy Balance methods are used. The Zanstra temperatures are estimated based on H I, He I, and He II RLs. Since the radiation-bounded condition is assumed in the Zanstra method, the Zanstra temperature estimated using the line with the highest ionization potential is most reliable. The Zanstra temperatures are used with the preferential order of He II, He I, and H I. A typical uncertainty in the Zanstra temperature is as large as 10,000 K in the worst case (Phillips 2003). Preite-Martinez et al. (1991) reported that the typical uncertainty in the temperature with the Energy Balance

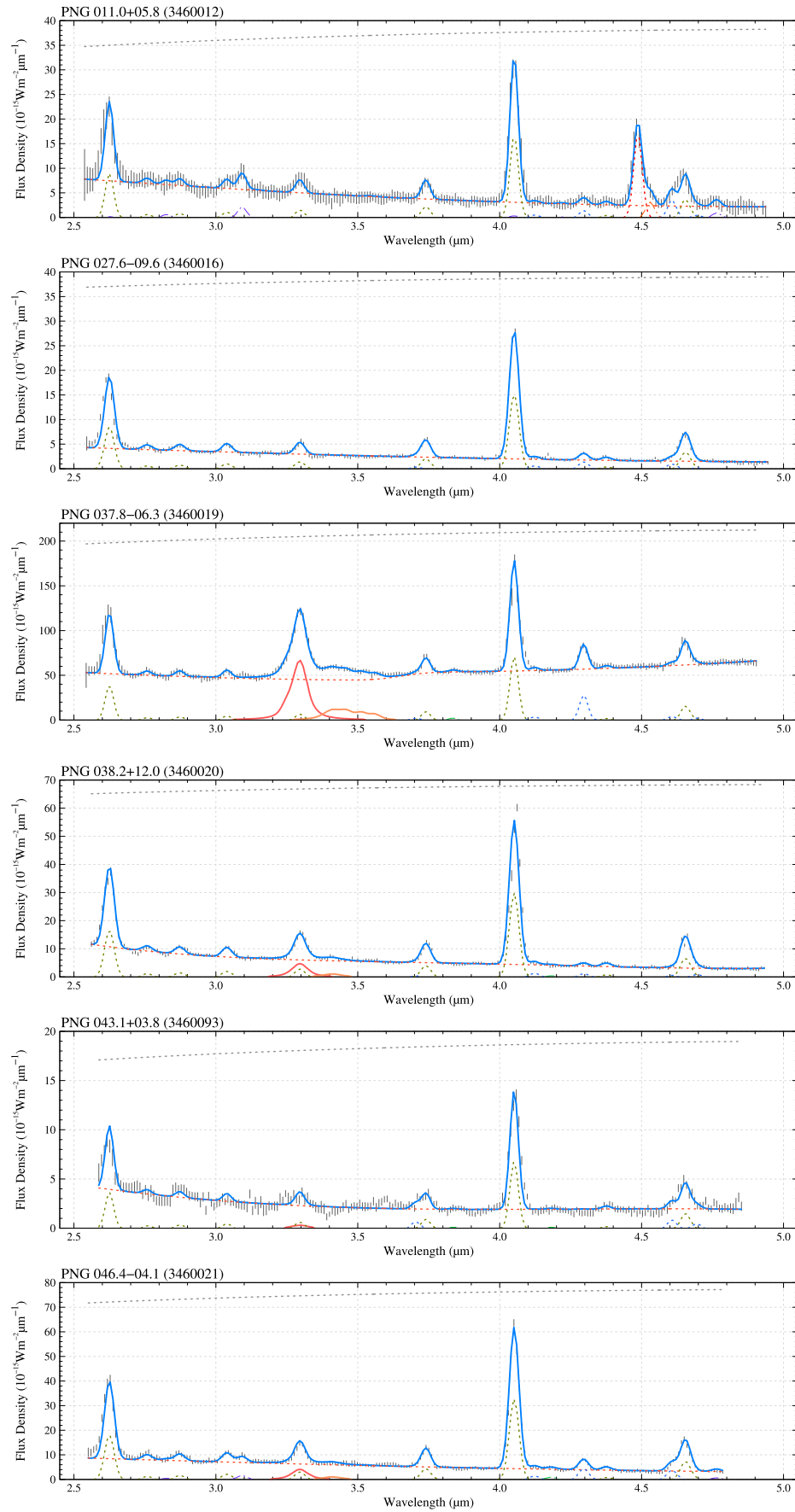


Figure 7. (Continued.)



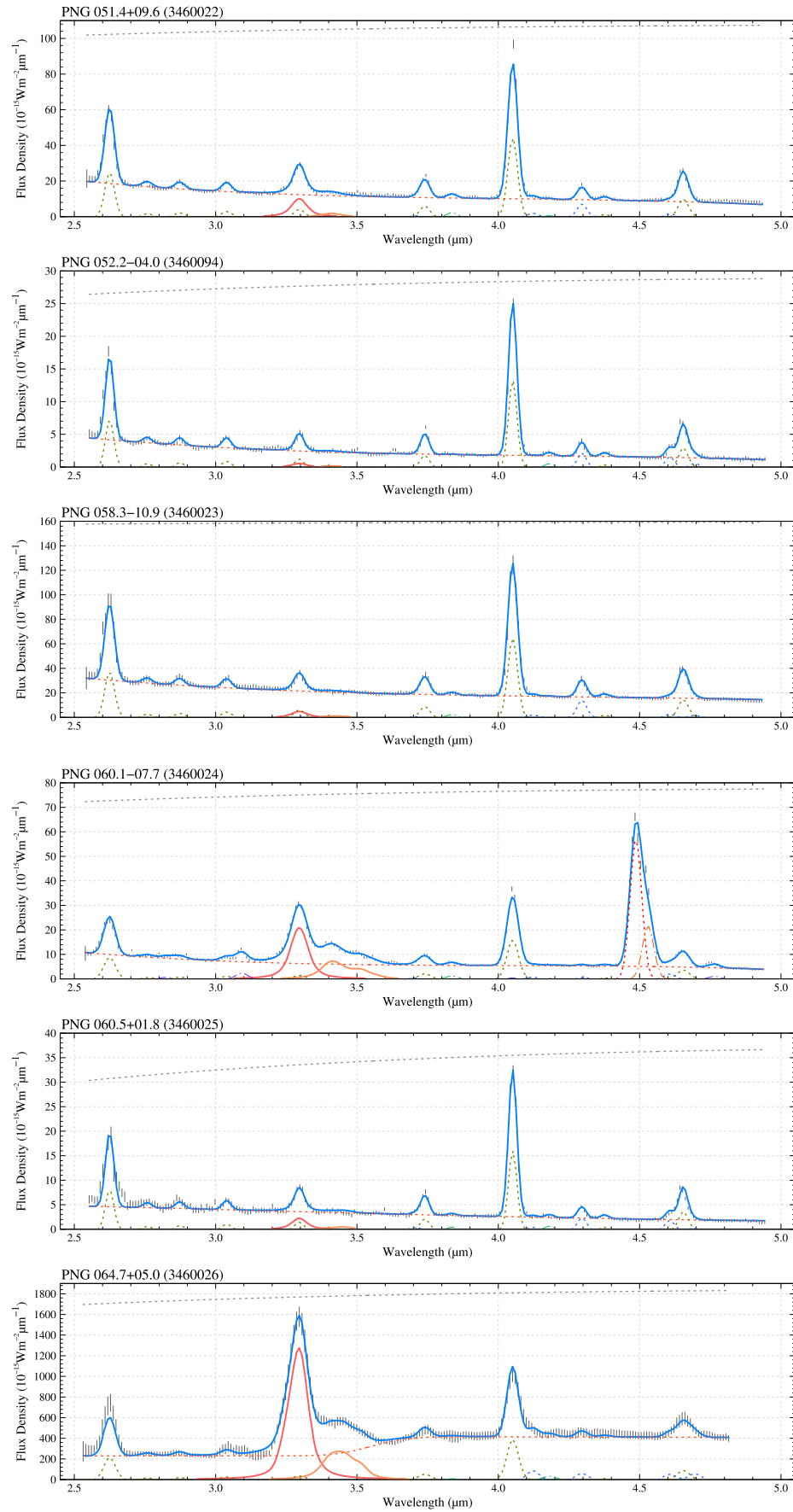


Figure 7. (Continued.)

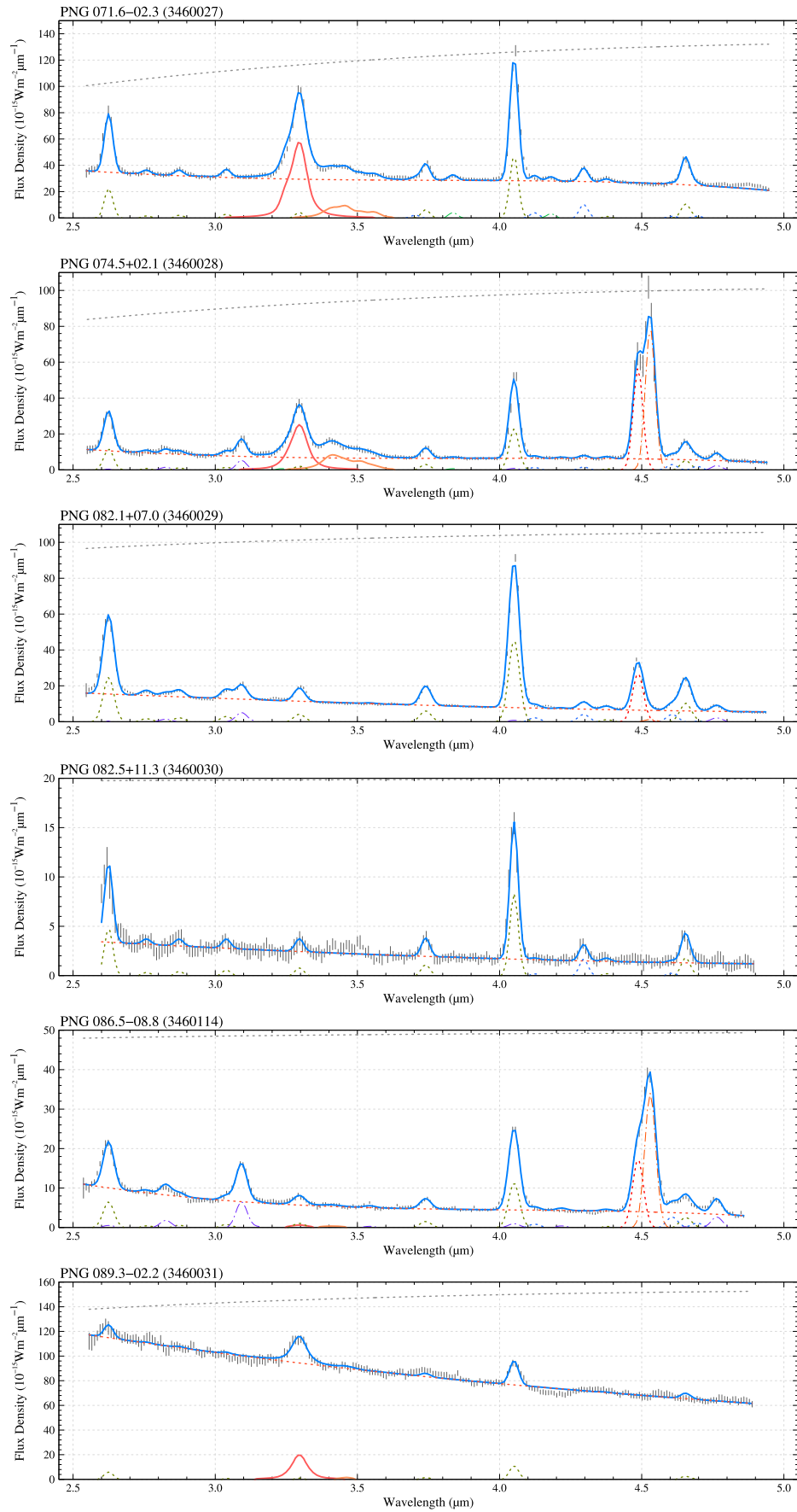


Figure 7. (Continued.)

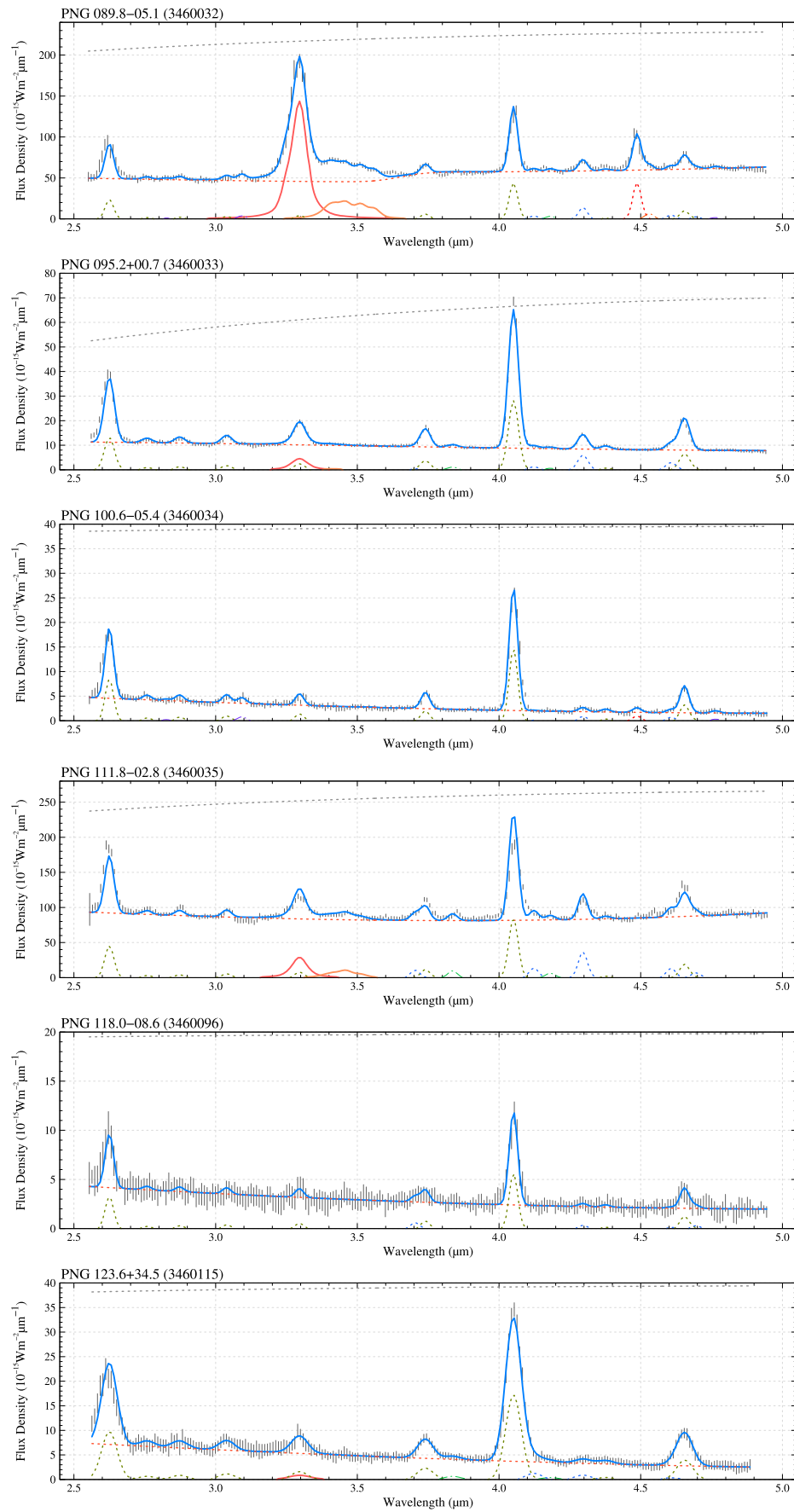


Figure 7. (Continued.)

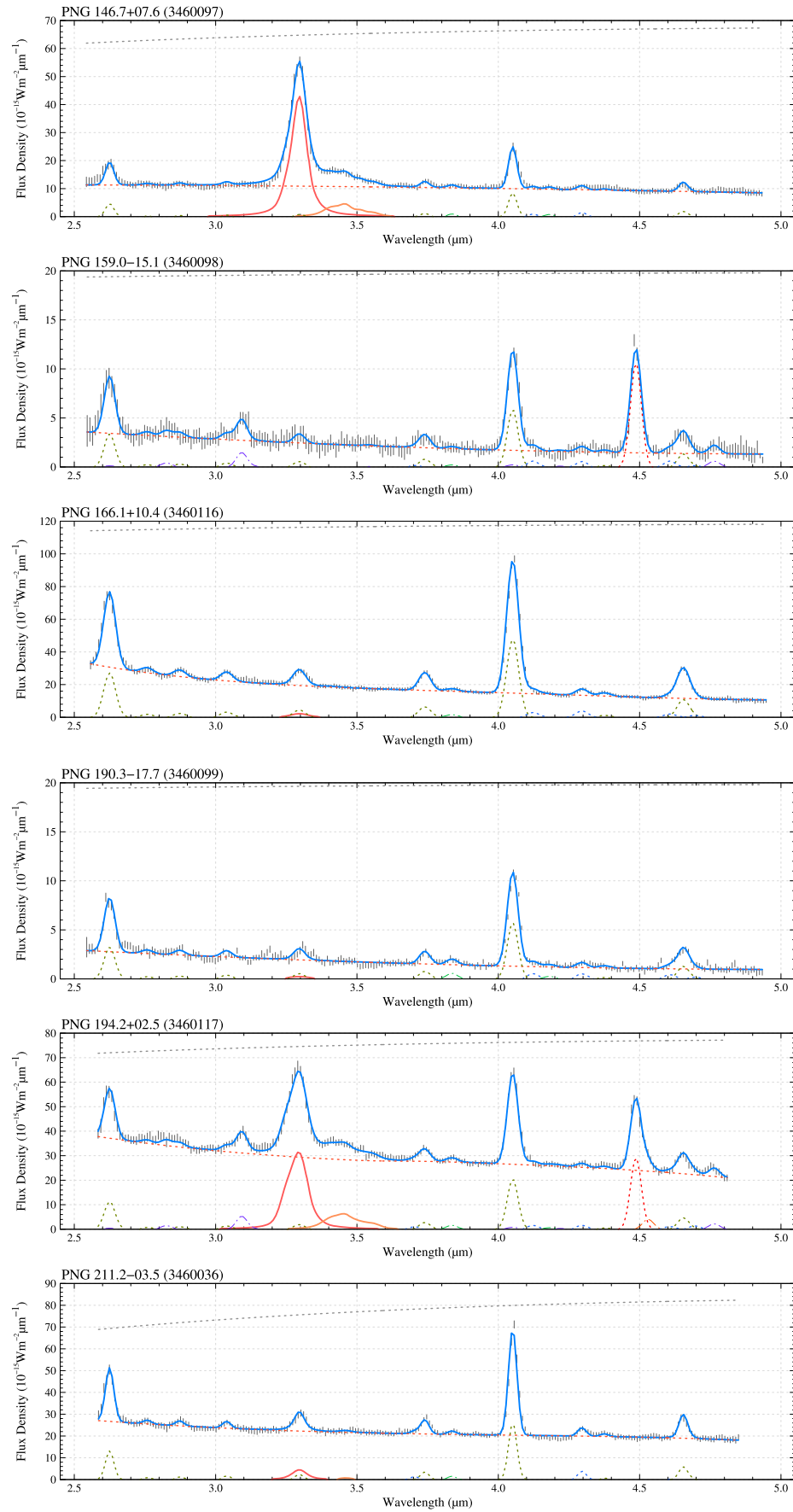


Figure 7. (Continued.)



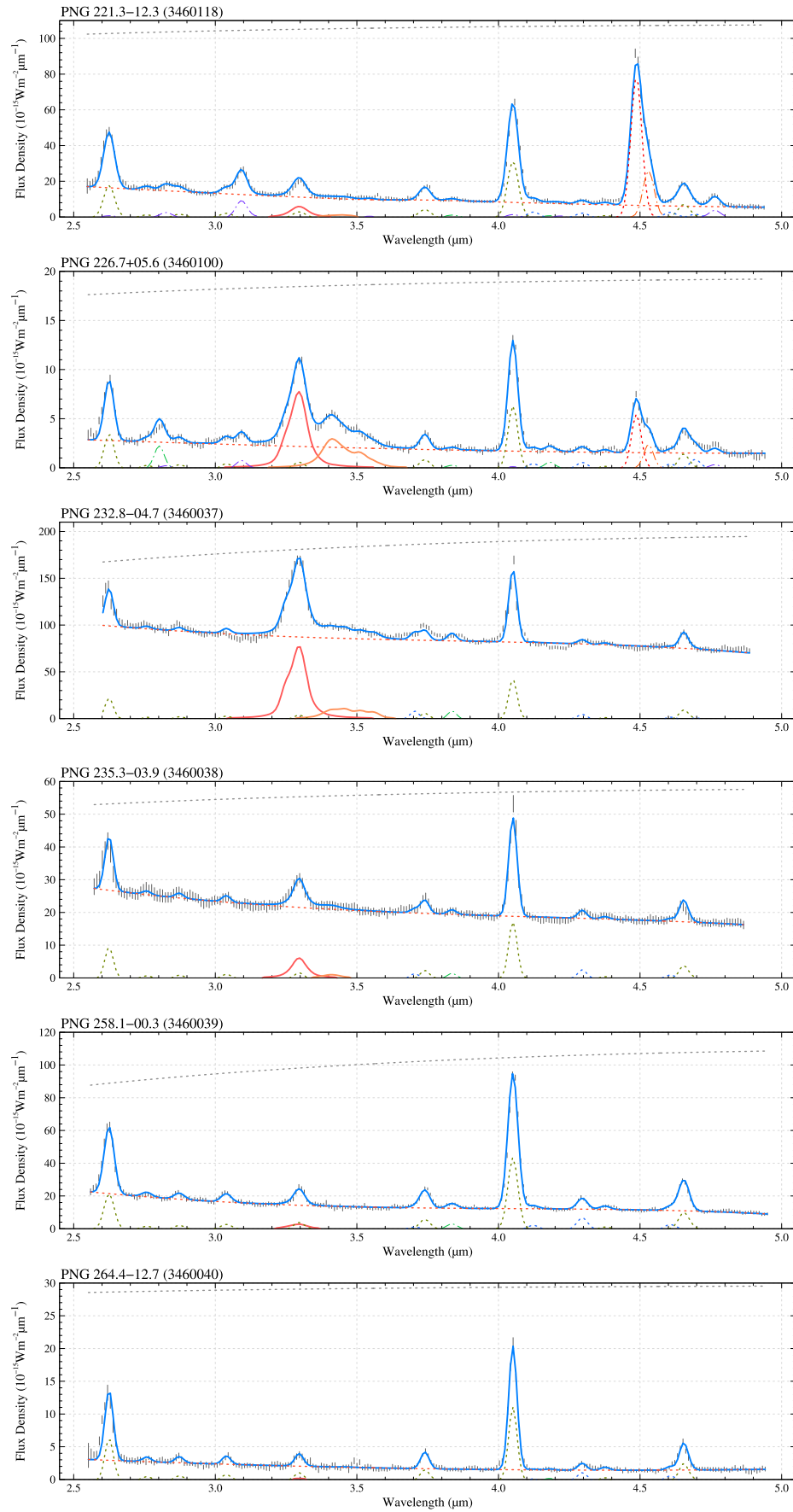


Figure 7. (Continued.)

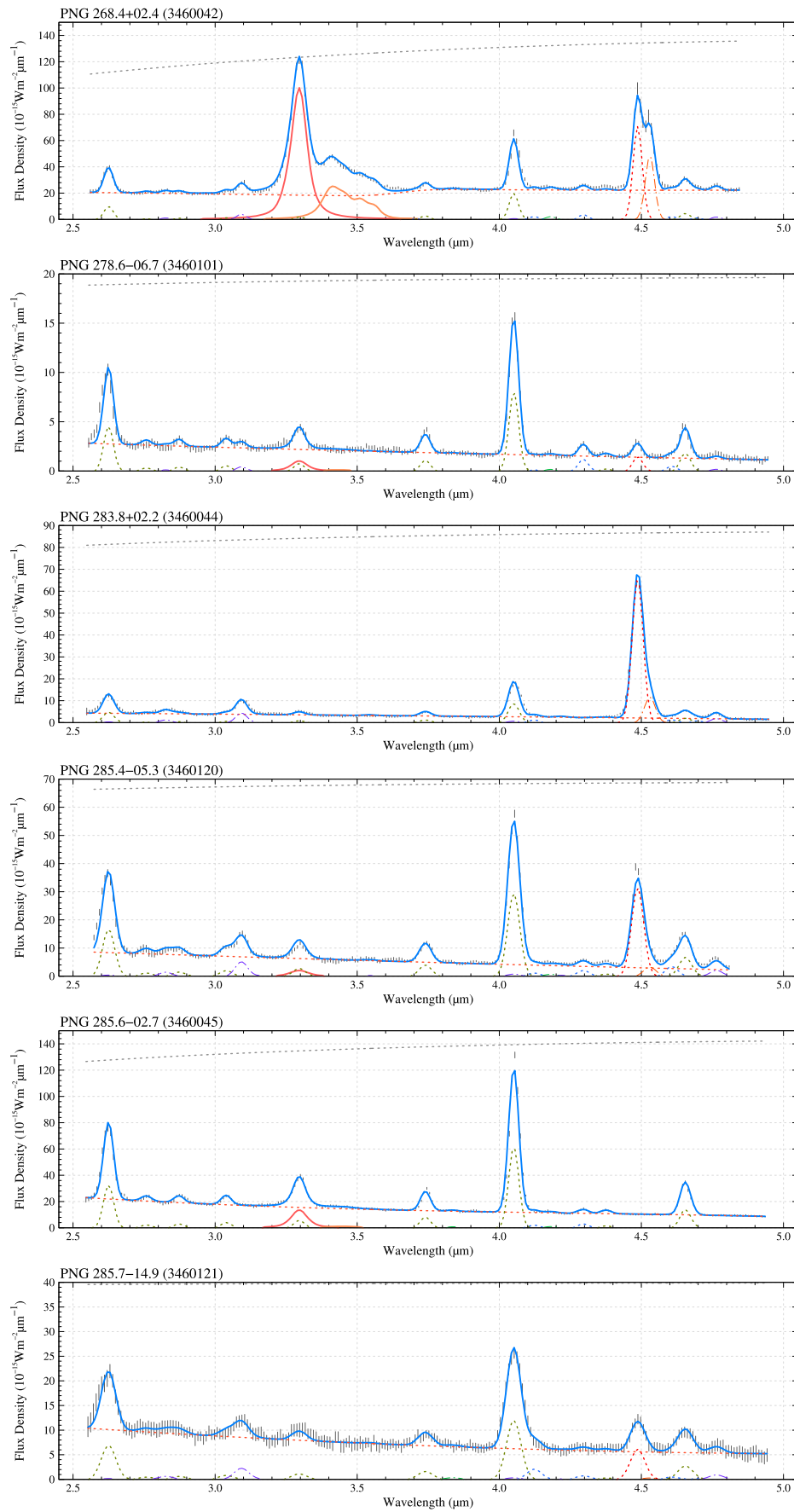


Figure 7. (Continued.)

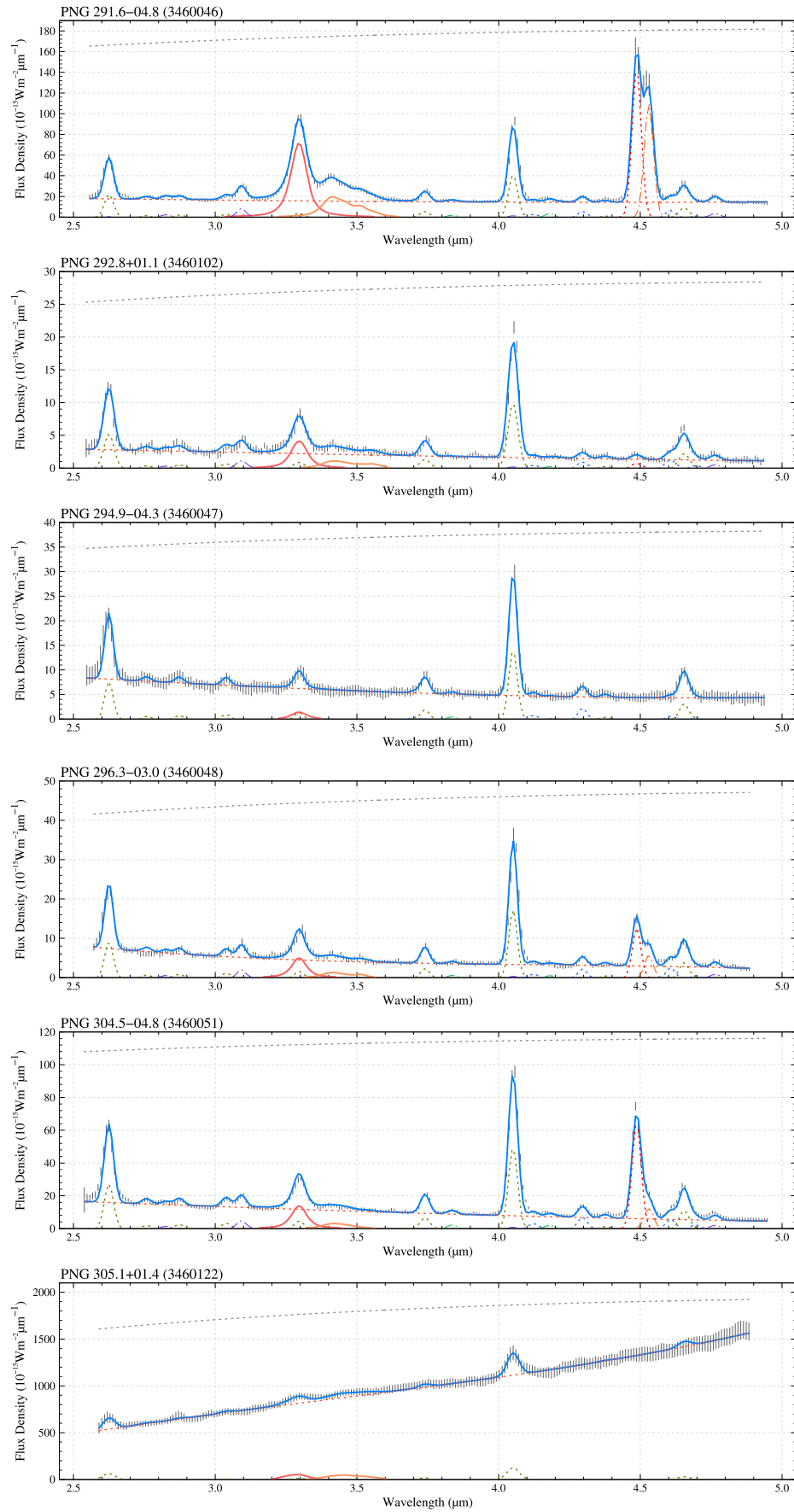


Figure 7. (Continued.)

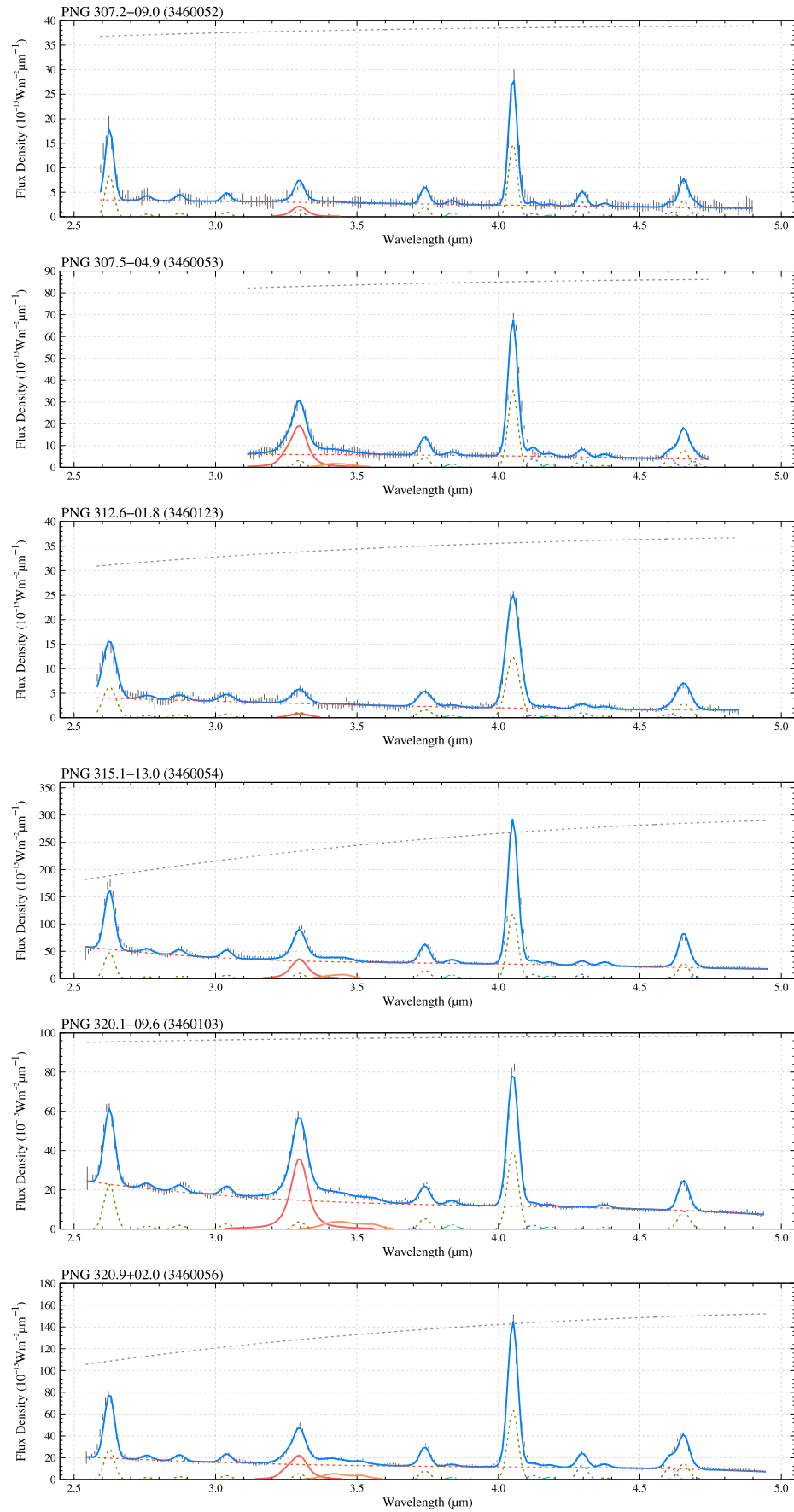


Figure 7. (Continued.)



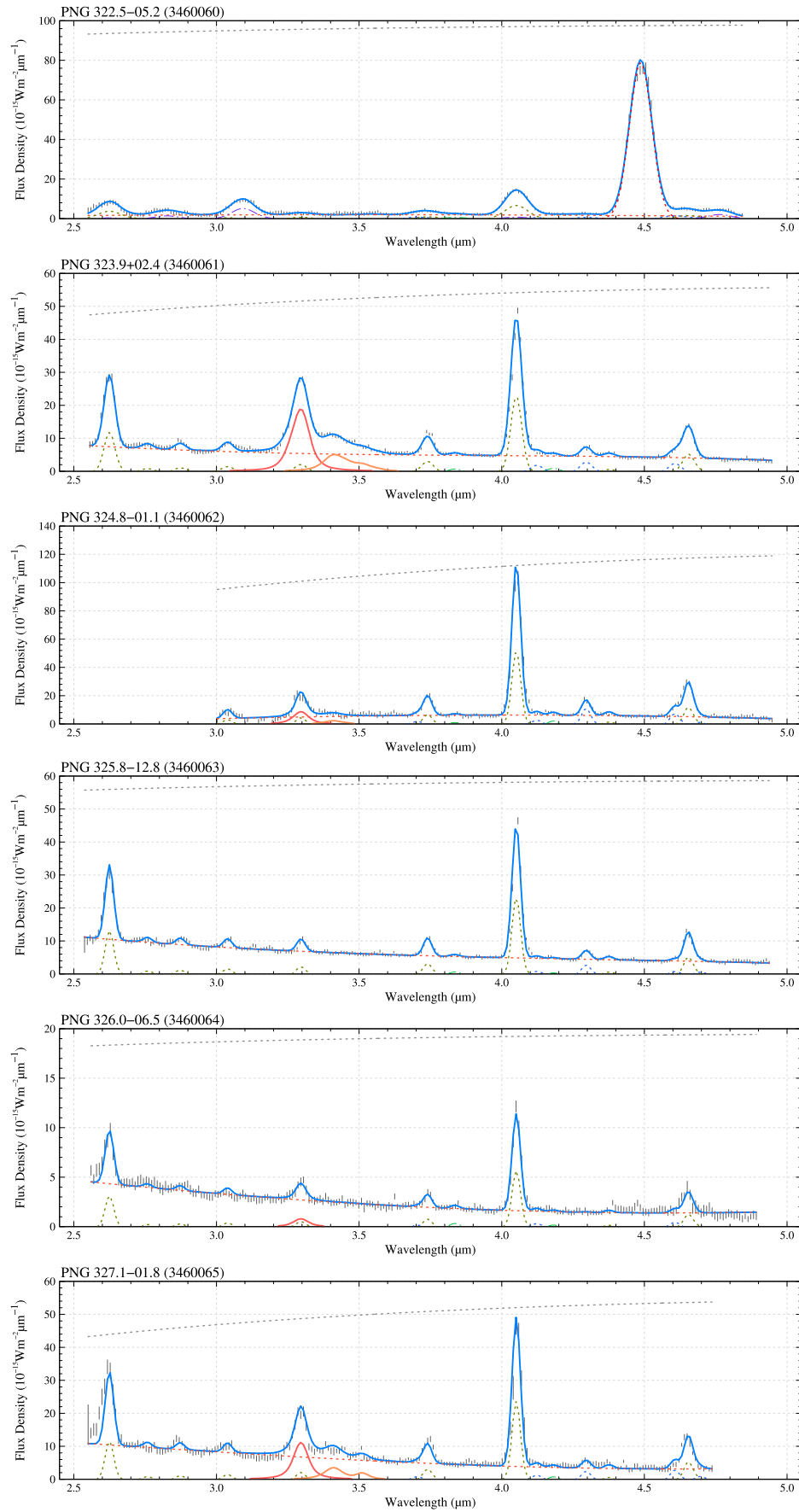


Figure 7. (Continued.)

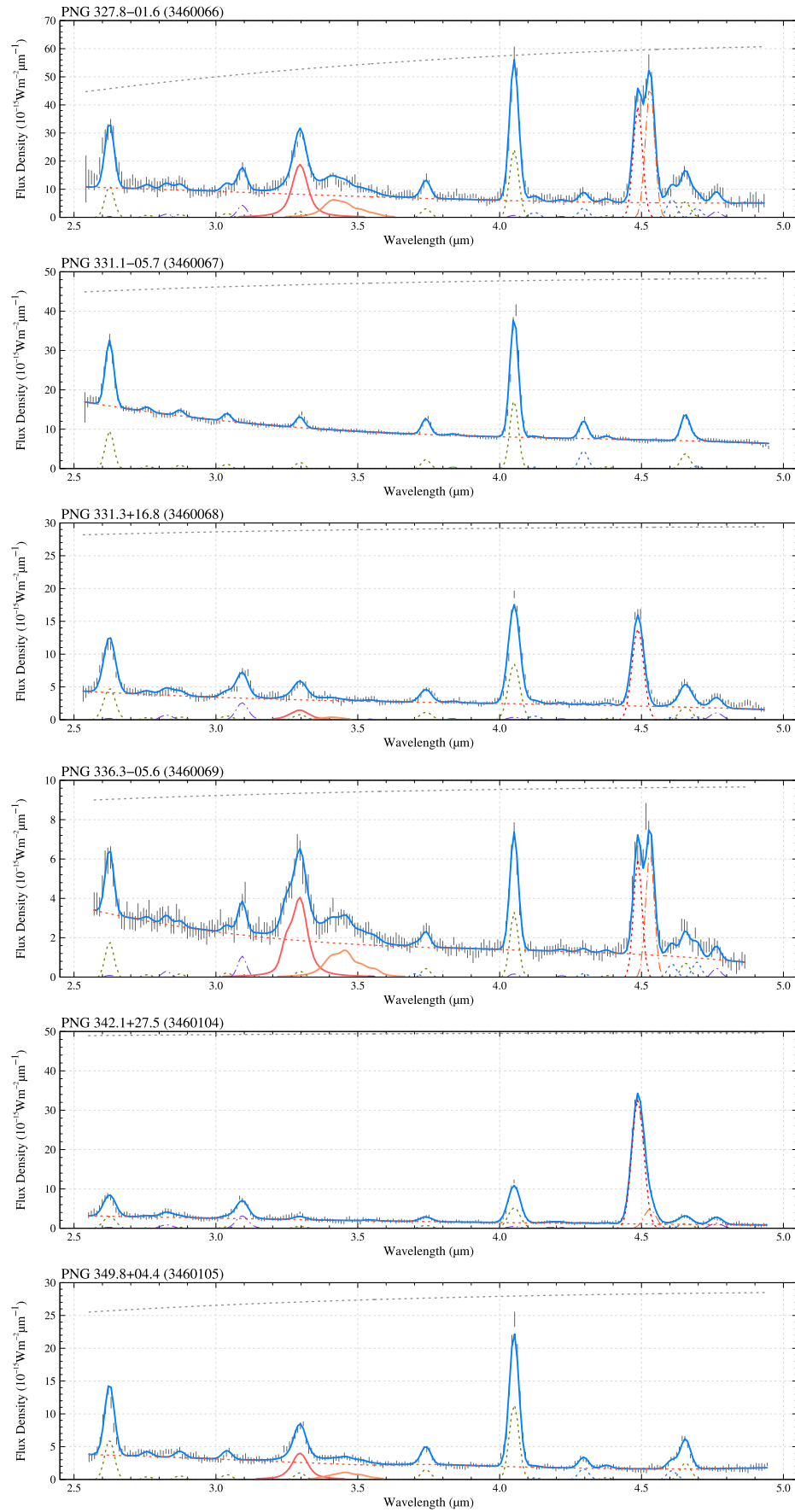


Figure 7. (Continued.)

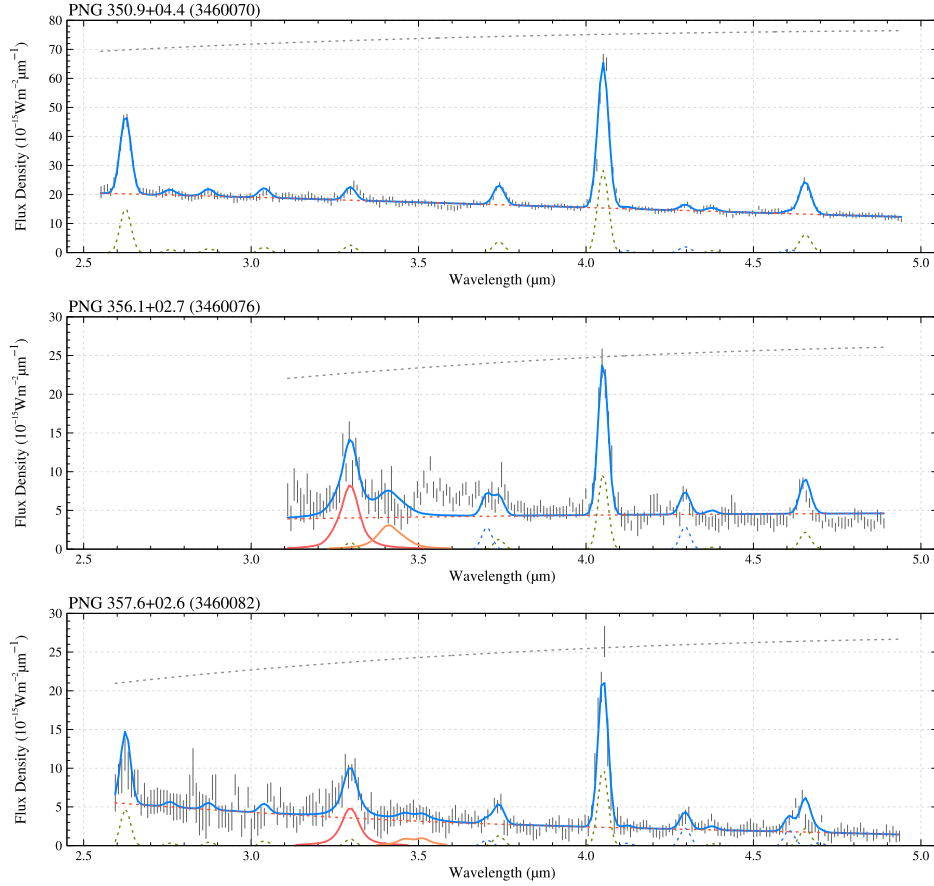
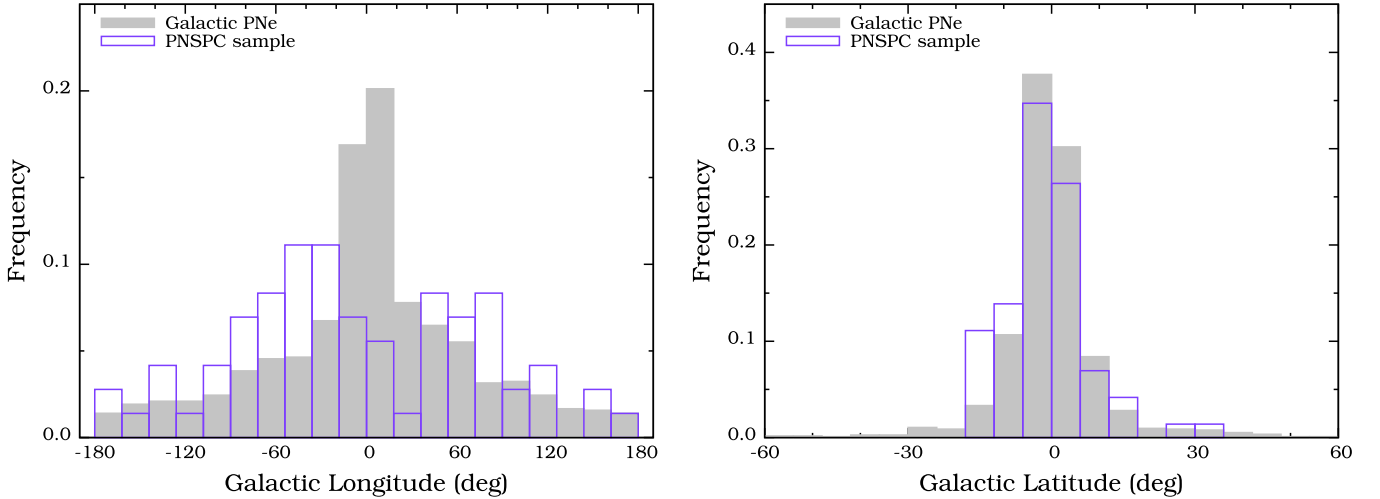


Figure 7. (Continued.)

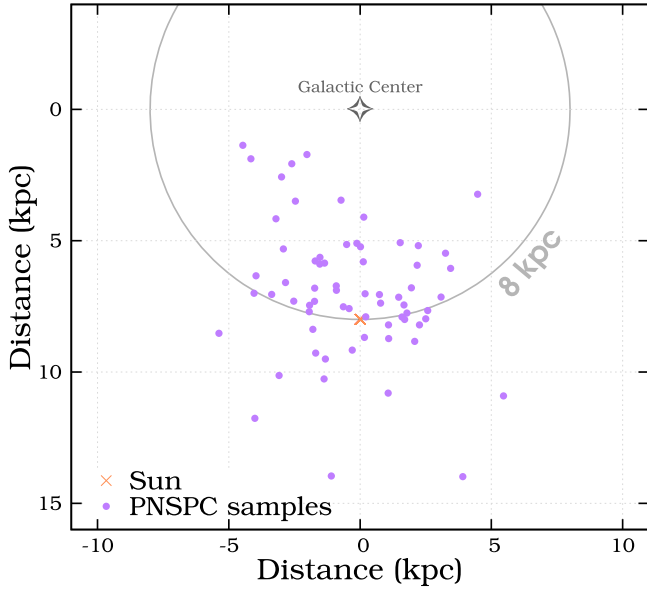


**Figure 8.** Histograms of PNe in the Galactic coordinates. The left panel shows the histogram of the Galactic longitude, while the right panel shows that of the Galactic latitude. The histograms for the PNSPC objects are shown by the blue boxes. The gray filled boxes represent the distribution of all the Galactic PNe in Acker et al. (1992).

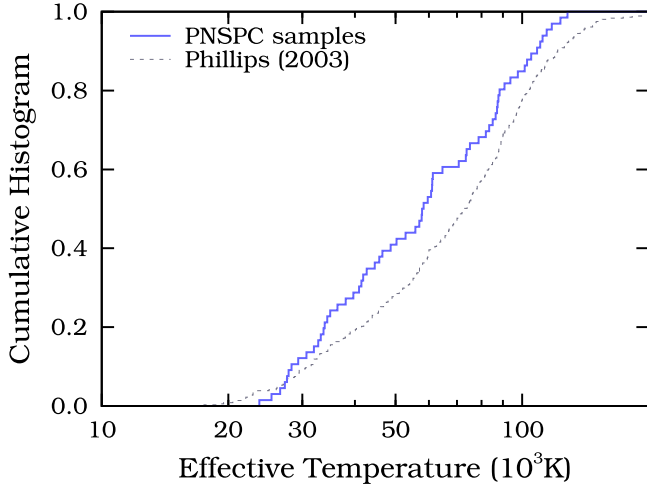
method is no more than 20%. We assume that a typical uncertainty in these methods is about 20%. Finally, the effective temperature was obtained for 67 of 72 objects in the catalog. When multiple references are available, the averaged value is adopted.

Figure 10 shows the distribution of the effective temperature in the PNSPC samples. Phillips (2003) collected the effective

temperature of 373 Galactic PNe using the Zanstra method. Their samples were widely collected from literature without any limitation. Thus, we assume that their samples are not biased and that the temperature distribution of Phillips' samples represents that of the whole Galactic PNe (gray dashed line). The median of the effective temperature for the PNSPC samples is about 50,000 K, while that for Phillips' data is about



**Figure 9.** Distribution of the PNSPC samples on the Galactic Plane. The red cross shows the locus of the Sun. The purple circles show the positions of the PNSPC samples.

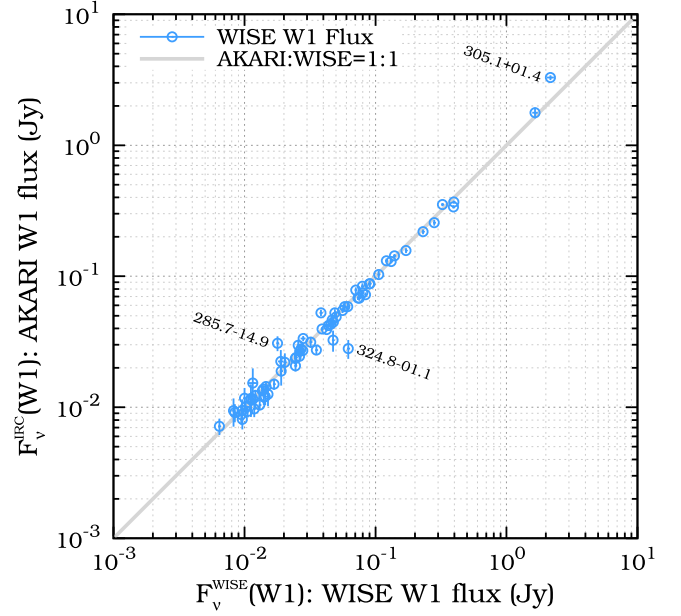


**Figure 10.** Cumulative histogram of the effective temperature of Galactic PNe. The blue solid line shows the histogram for the PNSPC samples. The gray dashed line represents the temperature distribution of all the Galactic PNe in Phillips (2003).

80,000 K. Figure 10 indicates that the PNSPC samples are biased toward low-temperature or young PNe, possibly due to the target selection bias: the targets are limited by the apparent size ( $\sim 3''$ ).

### 5.2. Accuracy of the Absolute Intensity

One of the major characteristics of the PNSPC catalog is slitless spectroscopy in the Np-window, which allows us to collect all of the flux from the target. Since the IRC spectrum totally covers the *WISE* W1 band, the accuracy of the absolute flux density is evaluated by comparing the PNSPC spectra with the *WISE* W1 photometry from *WISE* All-Sky Release Catalog (Cutri et al. 2012). Define  $R^{W1}$  as the relative response curve of the *WISE* W1 band (per photon; Wright et al. 2010). The flux in



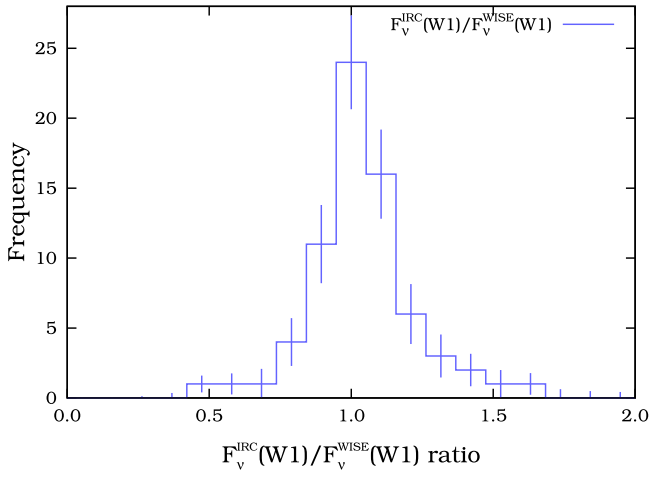
**Figure 11.** Flux comparison between the *AKARI* and *WISE* observations. The vertical axis shows the flux in the *WISE* W1 filter estimated from the *AKARI*/IRC spectra, while the horizontal axis shows the flux in the *WISE* All-Sky Release Catalog (Cutri et al. 2012). When the deviation between the *AKARI* and *WISE* W1 fluxes is larger than 50%, the data point is labeled with the PNG ID.

the W1 band estimated from the IRC spectrum is defined as

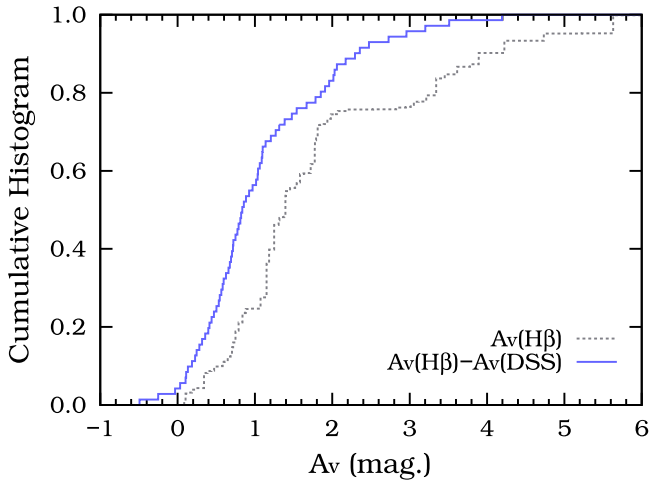
$$F_{\nu}^{\text{IRC}}(\text{W1}) = \frac{\int F_{\nu}^{\text{obj}}(\nu) R^{\text{W1}}(\nu) \frac{d\nu}{\nu}}{\int \left( \frac{\nu_{\text{iso}}}{\nu} \right)^2 R^{\text{W1}}(\nu) \frac{d\nu}{\nu}}, \quad (10)$$

where  $\nu_{\text{iso}}$  is the isophotal frequency of the W1 band ( $3.35 \mu\text{m}$ ) and  $F_{\nu}^{\text{obj}}$  is the observed flux density in units of Jy. Figure 11 shows the estimated flux density in the W1 band against the *WISE* W1 catalog data. It confirms that the W1 flux density estimated from the IRC spectra,  $F_{\nu}^{\text{IRC}}(\text{W1})$ , is consistent with the *WISE* W1 data in a wide range from about 0.01 to 5 Jy. The histogram of the *AKARI* to *WISE* W1 flux ratio is shown in Figure 12. Errors in Figure 12 are calculated by a Monte Carlo simulation with the uncertainty in the flux ratio. The ratios are distributed with the mean of  $1.01 \pm 0.01$  and the standard deviation of  $0.21 \pm 0.02$ . Three PNe (PNG 285.7–14.9, 324.8–01.1, and 305.1+01.4) show significant deviations larger than 50%. For PNG 285.7–14.9 and 305.1+01.4, their *AKARI* W1 fluxes are larger than those of *WISE*. Both objects are apparently extended in *WISE* W1 images. The deviations may be attributed to the loss of the flux in *WISE* photometry. Furthermore, PNG 305.1+01.4 is the brightest object in the PNSPC sample. The *WISE* detector starts to saturate in the W1 band at  $\sim 0.17$  Jy (Wright et al. 2010). This may also cause an error in the *WISE* photometry of PNG 305.1+01.4. The *AKARI* W1 flux of PNG 324.8–01.1 is smaller than the *WISE* W1 flux. PNG 324.8–01.1 is located in a crowded region, so that the IRC field of view may not be wide enough to estimate background emission accurately. The loss of the *AKARI* W1 flux is attributable to the error in the estimation of the





**Figure 12.** Distribution of the flux ratio: the flux in the *WISE* W1 filter estimated from the *AKARI*/IRC spectra to that from the *WISE* All-Sky Release Catalog.

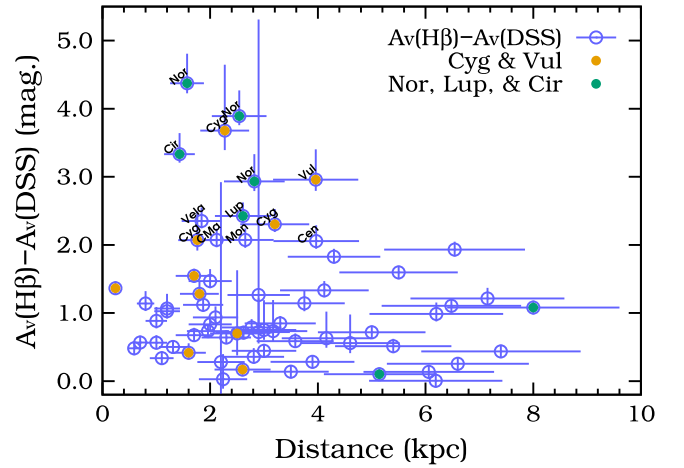


**Figure 13.** Cumulative histogram of the extinction at the V-band. The dotted line shows the histogram of the total extinction toward the PN. The solid line shows the histogram of the extinction without the contribution from the interstellar medium.

background emission. Thus, although the absolute continuum intensity of PNG 324.8–01.1 may be erroneous, the intensity of emission features is not affected by the estimation of background emission. The absolute flux of the present catalog is concluded to be accurate within 20%.

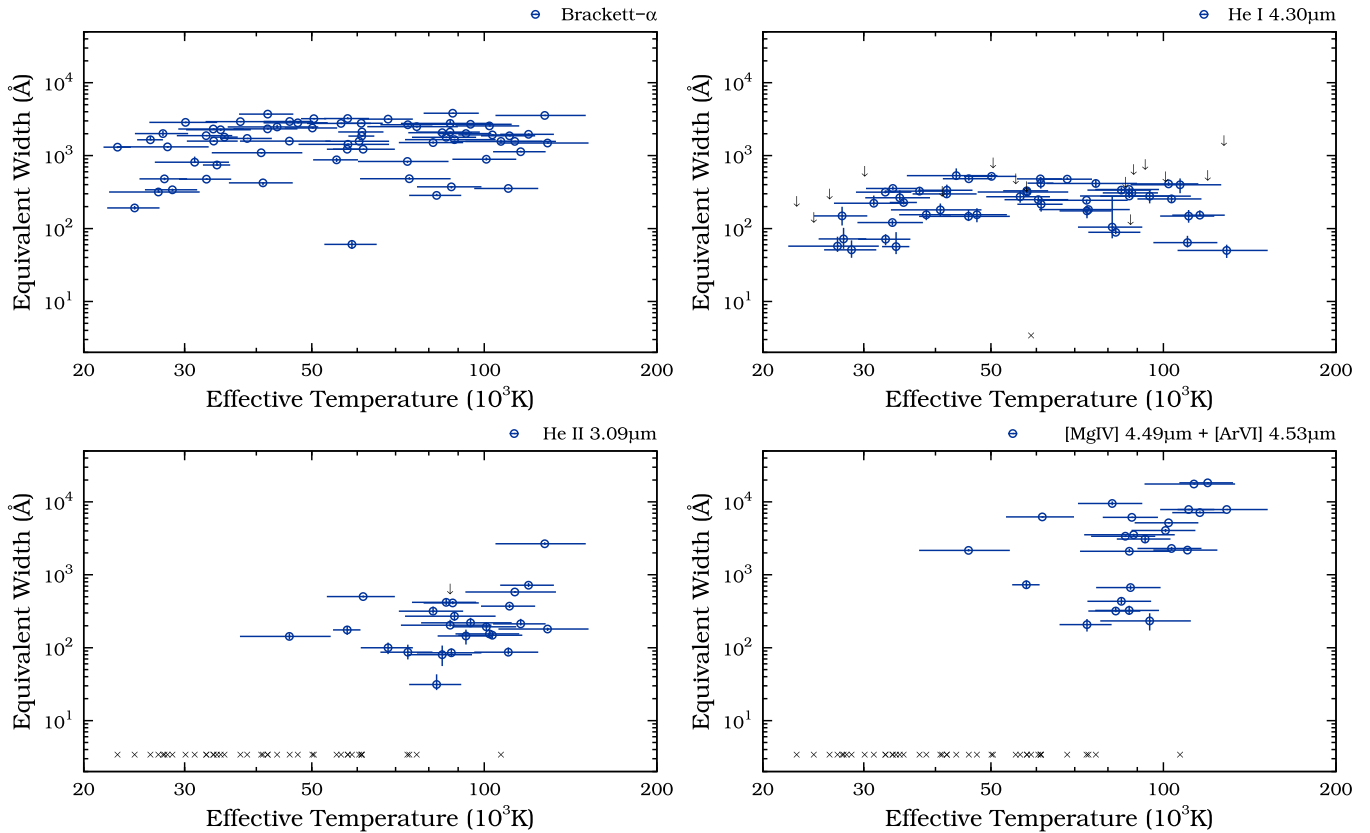
### 5.3. Extinction within the Individual Nebula

The PNSPC spectra include Brackett- $\alpha$  (Br $\alpha$ ), the hydrogen RL at  $4.05 \mu\text{m}$ , which enables us to estimate the extinction toward the objects using the relative intensity of H $\beta$  to Br $\alpha$  (the intensity of H $\beta$  is from Acker et al. 1992; see Section 4.2). The cumulative histogram of the total extinction toward the PNSPC samples,  $A_V(\text{H}\beta)$ , is shown by the dotted line in Figure 13. The extinction toward the PNSPC samples are less than 6 mag. at the V-band. About 70% of the PNSPC samples have  $A_V(\text{H}\beta) < 2 \text{ mag.}$ , suggesting that most of the PNSPC samples are not heavily obscured. This is consistent with the Galactic distribution of the PNSPC sample shown in Section 5.1.1.



**Figure 14.** Net extinction at the V-band is plotted against a distance toward the PN. Errors in the distance are assumed to be about 20%. The name of the constellation is added next to the symbols for the PN with the net extinction  $> 2.0 \text{ mag.}$  The PNe in Cygnus and Vulpecula are indicated by the orange filled circles, while the PNe in Norma, Lupus and Circinus are indicated by the green filled circles.

The extinction values calculated in Section 4.2 include the contribution from their circumstellar envelope,  $A_{V,\text{PN}}$ , and the interstellar medium to the object,  $A_{V,\text{ISM}}$ . The contribution from the interstellar medium,  $A_{V,\text{ISM}}$ , can be estimated using the extinction map based on the Digitized Sky Survey (Dobashi et al. 2005a, 2005b). The extinction map is given in the resolution of  $2' \times 2'$ . We calculate the averaged extinction values within a  $6'$  radius around the object  $A_V$  (DSS) and assumed them as  $A_{V,\text{ISM}}$ . Dobashi et al. (2005a) showed that a systematic error in  $A_V$  (DSS) is typically less than 0.2 mag for  $A_V$  (DSS)  $< 5 \text{ mag.}$  We estimate the uncertainty in  $A_{V,\text{ISM}}$  from the scatter of  $A_V$  (DSS) within the  $6'$  circle taking into account the systematic error. Note that the extinction  $A_{V,\text{ISM}}$  estimated here may have a large uncertainty. The extinction from the circumstellar envelope, defined by  $A_V(\text{H}\beta) - A_V(\text{DSS})$ , is shown by the solid line in Figure 13. Some objects show negative extinction values ( $\sim -0.5 \text{ mag.}$ ), possibly due to the overestimation of  $A_{V,\text{ISM}}$ . Figure 14 shows the dependence of  $A_V(\text{H}\beta) - A_V(\text{DSS})$  on the distance toward the PN. We find that PNe with extinction  $> 2 \text{ mag}$  frequently appear around 2 kpc. The name of the constellation is added next to those PNe, indicating that most of them are in the regions around Cygnus and Norma as indicated by the orange and green symbols. The extinction  $A_{V,\text{ISM}}$  for these PNe may be underestimated or have large local fluctuations, but this affects only about 10% of the PNSPC samples. Except for the objects in the Cygnus and Norma regions, Figure 14 does not show any trend with the distance, suggesting that the  $A_V(\text{H}\beta) - A_V(\text{DSS})$  value is attributable to the extinction internal to the PN. Figure 13 indicates that the median of the circumstellar extinction is about 0.8 mag. and about 40% of the PNe show extinction larger than 1 mag. suggesting that the circumstellar envelope of a PN is generally optically thick in the UV. Ohsawa et al. (2012) investigated the spatial distribution of mid-infrared dust emission of PNG 095.2+00.7, a young and UV-thick Galactic PN, and suggest that the evolution of a UV-thick circumstellar envelope may have a significant impact on the appearance of the mid-infrared spectrum of the PN. The present results suggest that PNe spectra should be investigated



**Figure 15.** Equivalent widths of Brackett- $\alpha$  at  $4.05 \mu\text{m}$  (upper left), He I at  $4.30 \mu\text{m}$  (upper right), He II at  $3.09 \mu\text{m}$  (lower left), and [Mg IV] at  $4.49 \mu\text{m}$  + [Ar VI] at  $4.53 \mu\text{m}$  (lower right) against the effective temperature. The downward arrows indicate a  $3\sigma$  upper limit. When the emission is not detected, the effective temperature is indicated by the crosses.

by taking into account the radiation transfer within the nebula and the gradient in the dust temperature.

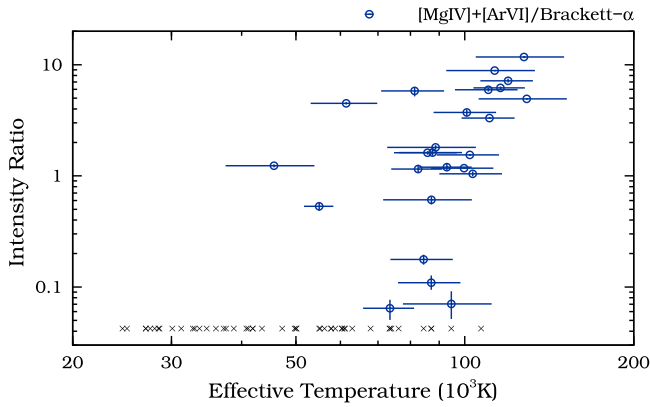
#### 5.4. Equivalent Width of Ionized Gas Emission Lines

The spectra in the PNSPC catalog contain several emission lines, including hydrogen and helium RLs and fine-structure lines of [Mg IV] at  $4.49 \mu\text{m}$  and [Ar VI] at  $4.53 \mu\text{m}$ . These emission lines should reflect the characteristics of the radiation field of the circumstellar envelope. Figure 15 shows the equivalent widths of Br $\alpha$ , He I at  $4.30 \mu\text{m}$ , He II at  $3.09 \mu\text{m}$ , and the summation of [Mg IV] and [Ar VI] against the effective temperature of the central star. Since the [Mg IV] and [Ar VI] lines are close to each other, the equivalent width of the summation of these lines is used. The data whose S/N of the equivalent width is less than three are replaced by  $3\sigma$  upper limits. When the emission is not detected, the effective temperature is indicated by the crosses. The equivalent width of Br $\alpha$  is typically about  $(1\text{--}3) \times 10^3 \text{ \AA}$  and does not show a clear dependence on the effective temperature. The He I line is seen from 20,000 to 150,000 K. Its equivalent width is almost constant at about  $2 \times 10^2 \text{ \AA}$ . It shows a small increase from 30,000 to 50,000 K, possibly explained by the increase in photons with enough energy to ionize He with increasing effective temperature. When the effective temperature exceeds 50,000 K, the He I equivalent width stops increasing and shows a small decrease. At the same time, the He II lines start to be detected, suggesting that high-energy photons are in part used to ionize He $^+$ . The decrease in the He I equivalent width is attributable to the ionization balance between He $^+$  and He $^{++}$ .

The He II, [Mg IV], and [Ar VI] lines are not detected until the effective temperature exceeds about 50,000 K. The ionization potentials of He II, Mg IV, and Ar VI are 54.41, 109.27, and 91.00 eV. Although the ionization potential of He II is about two times lower than those of Mg IV, and Ar VI, the He II, [Mg IV], and [Ar VI] lines start to appear at almost the same temperature in Figure 15. The equivalent width of the He II line is typically about  $2 \times 10^2 \text{ \AA}$ , not showing a clear trend with the effective temperature. The equivalent width of the [Mg IV] + [Ar VI] becomes as large as  $8 \times 10^3 \text{ \AA}$  at a high temperature of about 100,000 K. Figure 16 shows the intensity ratio of the [Mg IV] + [Ar VI] to Br $\alpha$ . The effective temperature of the PNe with the ratio larger than unity ranges about 500,000–150,000 K. The [Mg IV] and [Ar VI] lines are an indicator of high-temperature PNe as expected. But note that some PNe with high effective temperature show a ratio less than unity. A ratio less than unity does not necessarily indicate a PN with a low effective temperature.

#### 5.5. PAH Features of Galactic PNe in the Near-infrared

The PAH features appear both in the near- and mid-infrared, and the PAH features in the mid-infrared are stronger than those in the near-infrared (e.g., Schutte et al. 1993; Draine & Li 2001). The PAH features of PNe in the mid-infrared have been intensively investigated with *Spitzer* (e.g., Stanghellini et al. 2007, 2012; Bernard-Salas et al. 2009). Surveys of the  $3.3 \mu\text{m}$  PAH feature in Galactic PNe were carried out by Roche et al. (1996) and Smith & McLean (2008). Their samples were, however, biased toward carbon-rich objects. The PNSPC

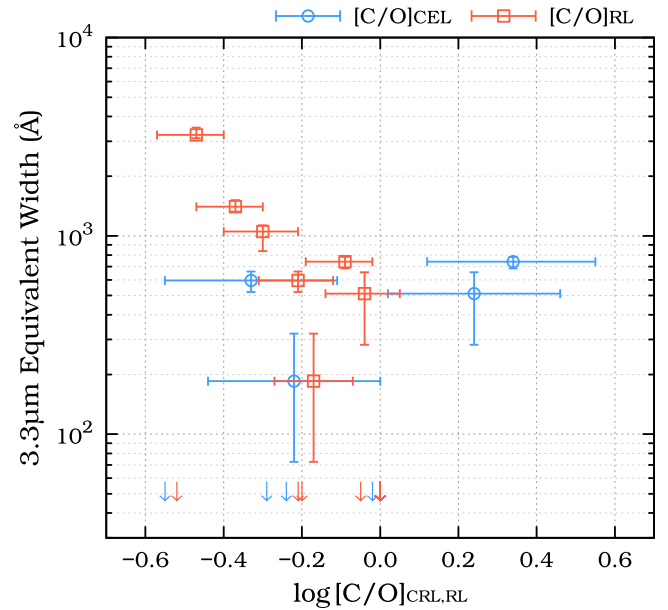


**Figure 16.** Intensity ratio of the summation of [Mg IV] and [Ar VI] to Brackett- $\alpha$  against the effective temperature. When the emission of [Mg IV] and [Ar VI] is not detected, the effective temperature is indicated by the crosses.

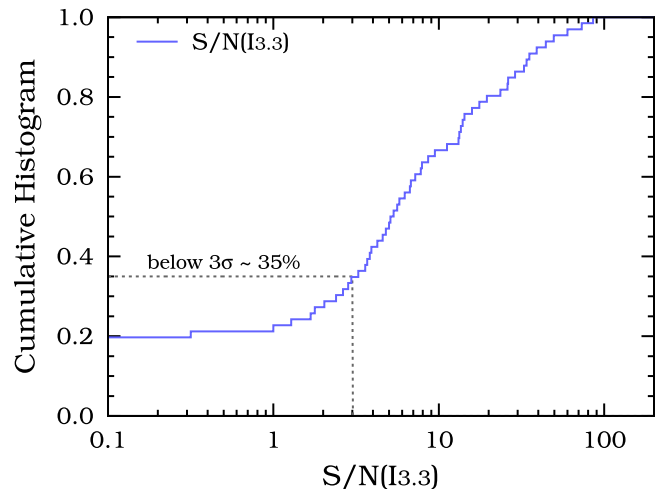
samples are not selected based on their chemistry. The PNSPC catalog provides a useful data set to investigate the near-infrared PAH features in Galactic PNe.

The type of dust features (carbon-rich or oxygen-rich) is thought to be closely related to the abundance ratio of nebular gas. Several studies have suggested that the intensities of the PAH features increases with the carbon-to-oxygen abundance (C/O) ratio (e.g., Cohen et al. 1986; Roche et al. 1996; Cohen & Barlow 2005; Smith & McLean 2008). However, PAH formation both in carbon-rich and oxygen-rich environments has been suggested recently (Guzman-Ramirez et al. 2011, 2014). Delgado-Inglada & Rodríguez (2014) measured the C/O ratio of 51 Galactic PNe and confirmed that the PAH emission can be seen even in PNe with  $C/O < 1$ . The C/O ratios of 11 PNe in the PNSPC catalog were measured by Delgado-Inglada & Rodríguez (2014). Figure 17 shows the equivalent width of the  $3.3 \mu\text{m}$  PAH feature against the C/O ratio. The C/O ratio measured by collisionally excited lines (CEL) is indicated by the blue circles, while that measured by RLs is shown by the red squares. The downward arrows indicate non-detection of the  $3.3 \mu\text{m}$  PAH features. Figure 17 shows that the  $3.3 \mu\text{m}$  PAH emission is detected even in oxygen-rich environments, consistent with Delgado-Inglada & Rodríguez (2014). Due to the small number of samples, however, it is difficult to claim the correlation between the strength of the  $3.3 \mu\text{m}$  PAH feature and the C/O ratio. Extensive measurements of the C/O ratio are encouraged.

Figure 18 shows a cumulative histogram of the signal-to-noise ratio of the  $3.3 \mu\text{m}$  PAH feature (hereafter,  $S/N(I_{3.3})$ ). The gray dotted line shows the 35th percentile, which corresponds to  $S/N(I_{3.3}) = 3$ . The result suggests that about 65% of the PNSPC samples show band emission at  $3.3 \mu\text{m}$ . The  $5\text{--}38 \mu\text{m}$  spectra of Galactic PNe with *Spitzer*/IRS are available in Stanghellini et al. (2012). We counted the number of PNe with PAH features, obtaining 40% (60 of 150 PNe). Despite the intrinsic weakness of the near-infrared PAH features, the PAH detection rate is as high in the near-infrared as in the mid-infrared. The  $3.3 \mu\text{m}$  feature is thought to come from the smallest PAH populations, which are most fragile, and thus primarily reflect environment variations (Schutte et al. 1993; Allain et al. 1996a, 1996b; Micelotta et al. 2010). The PNSPC catalog is useful to investigate the evolution of the PAH features and the processing of PAHs in circumstellar environments.



**Figure 17.** Equivalent width of the  $3.3 \mu\text{m}$  PAH feature against the carbon-to-oxygen (C/O) ratio. The C/O ratios are from Delgado-Inglada & Rodríguez (2014). The C/O ratio measured by collisionally excited lines (CEL) is shown by the blue circles, while the C/O ratio measured by recombination lines (RL) is by the red squares. The arrows indicate the non-detection of the  $3.3 \mu\text{m}$  PAH feature.



**Figure 18.** Cumulative histogram of the signal-to-noise ratio of the  $3.3 \mu\text{m}$  PAH feature ( $S/N(I_{3.3})$ ). The gray dotted-line shows the 35th percentile, which is the non-detection rate of the  $3.3 \mu\text{m}$  PAH feature when the detection limit is defined as  $S/N(I_{3.3}) \geq 3$ .

## 6. SUMMARY AND CONCLUSION

Near-infrared ( $2.5\text{--}5.0 \mu\text{m}$ ) spectra of 72 Galactic PNe were obtained at high sensitivity using the *AKARI*/IRC. The PNSPC program provided a set of near-infrared spectra that collected all of the flux from the objects. The absolute flux of spectra provided by the PNSPC program is in agreement with the *WISE* W1 photometry within  $\sim 20\%$ . These spectra were compiled into the PNSPC catalog. The intensity and equivalent width of Brackett- $\alpha$ , He I at  $4.30 \mu\text{m}$ , He II at  $3.09 \mu\text{m}$ , the  $3.3 \mu\text{m}$  PAH feature, the  $3.4\text{--}3.5 \mu\text{m}$  aliphatic feature complex, [Mg IV] at  $4.49 \mu\text{m}$ , and [Ar VI] at  $4.53 \mu\text{m}$  were measured. The detection limit of the emission features was as faint as  $2 \times 10^{-16} \text{ W m}^{-2}$ . The PNSPC catalog is the largest data set of near-infrared spectra of Galactic PNe. It provides unique

information for the investigation of 2.5–5.0  $\mu\text{m}$  emission of Galactic PNe.

The Galactic coordinates of the PNSPC samples suggest that the objects in the PNSPC catalog are biased toward PNe in the Galactic disk rather than those in the Galactic bulge. The median of the effective temperature of the PNSPC samples is lower than that of the whole Galactic PNe by about 30,000 K. This indicates that the PNSPC catalog is biased toward young PNe. This bias possibly originates in the criterion on the apparent size in the target selection.

The PNSPC catalog also provides the extinction toward the objects. Even after the contribution from the interstellar medium is taken into account, Figure 13 suggests that 40% of PNe have extinction at the V-band larger than unity, which is attributed to the extinction in the circumstellar envelope. The present result suggests that for a large fraction of PNe a circumstellar envelope is optically thick in the UV.

The equivalent width of Brackett- $\alpha$  does not show a clear dependence on the effective temperature. The variations in the He I and He II equivalent widths can be attributed to the ionization balance between  $\text{He}^+$  and  $\text{He}^{++}$ . The [Mg IV] at 4.49  $\mu\text{m}$  and [Ar VI] at 4.53  $\mu\text{m}$  lines are only detected when the effective temperature becomes higher than 50,000 K. The [Mg IV] and [Ar VI] lines are good indicators of PNe with a hot central star.

The 3.3  $\mu\text{m}$  PAH feature is detected in about 65% of the PNSPC samples. This detection rate is comparable to that reported from mid-infrared observations (Stanghellini et al. 2012). The present result suggests that the AKARI/IRC is as sensitive as the *Spitzer*/IRS in terms of detecting the PAH features. The PNSPC catalog provides a suitable data set to investigate PAHs in circumstellar environments. The near-infrared PAH features are attributed to small-sized PAHs, which are sensitive to harsh environments and processed faster than larger ones. The processing and evolution of PAHs during the PN phase will be discussed in the forthcoming paper.

The present results are based on observations with AKARI, a JAXA project with the participation of ESA. We greatly appreciate all the people who worked in the operation and maintenance of those instruments. This publication makes use of data products from the *Wide-field Infrared Survey Explorer*, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. This work is supported in part by Grant-in-Aids for Scientific Research (25-8492, 23244021, and 26247074) by the Japan Society of Promotion of Science (JSPS).

## APPENDIX UNCERTAINTY IN $A_V$

There are several hydrogen RLs in the 2.5–5.0  $\mu\text{m}$  spectrum. Extinction can be estimated based on the intensity ratio of these lines. The accuracy of the estimated extinction heavily depends on the S/N of the intensity ratio. This chapter describes the method to derive the extinction based on the IRC spectrum and shows that a typical uncertainty in  $A_V$  is  $\sim 1$  mag for the PNSPC samples.

The observed intensity of the line emission  $X$  is given by  $I_X^{\text{obs}} = I_X e^{-\tau_X}$ , where  $I_X$  is the intrinsic intensity and  $\tau_X$  is the extinction at the line  $X$ . The observed intensity ratio of the line

$X$  to  $Y$  is

$$\frac{I_X^{\text{obs}}}{I_Y^{\text{obs}}} = \frac{I_X}{I_Y} \exp(-\alpha(X, Y)A_V), \quad (11)$$

where  $\alpha(X, Y)$  is defined by  $(\tau_X - \tau_Y)/A_V$ , which is a constant value specific to the extinction curve. Thus, the extinction is estimated by

$$A_V = \frac{\log x_{\text{int}} - \log x_{\text{obs}}}{\alpha(X, Y)}, \quad (12)$$

where  $x_{\text{int}}$  is  $I_X/I_Y$  and  $x_{\text{obs}}$  is  $I_X^{\text{obs}}/I_Y^{\text{obs}}$ . Given that the line  $X$  and  $Y$  are Brackett- $\alpha$  (Br $\alpha$ ) and  $\beta$  (Br $\beta$ ), respectively, the intrinsic intensity ratio is about 0.57 assuming the Case-B condition (Baker & Menzel 1938) with the electron density of  $10^4$  K and the electron density of  $10^4 \text{ cm}^{-3}$ . The ratio is less sensitive to either the electron temperature or density. We adopt the extinction curve given by Mathis (1990). The value of  $\alpha(\text{Br}\beta, \text{Br}\alpha)$  is about 0.041. The extinction estimated from the Brackett- $\beta$  to  $\alpha$  ratio is

$$A_V(\text{Br}\beta, \text{Br}\alpha) \simeq -24.6 \log x_{\text{obs}} - 13.7. \quad (13)$$

The uncertainty in  $A_V$  (Br $\beta$ , Br $\alpha$ ) is investigated. The measured intensity ratio ( $x_{\text{obs}}$ ) includes a statistical error. Define the error by  $\sigma$ . Assuming that the variation in  $x_{\text{obs}}$  is given by a normal distribution, the posterior probability distribution of the intensity ratio ( $\bar{x}_{\text{obs}}$ ) is given by

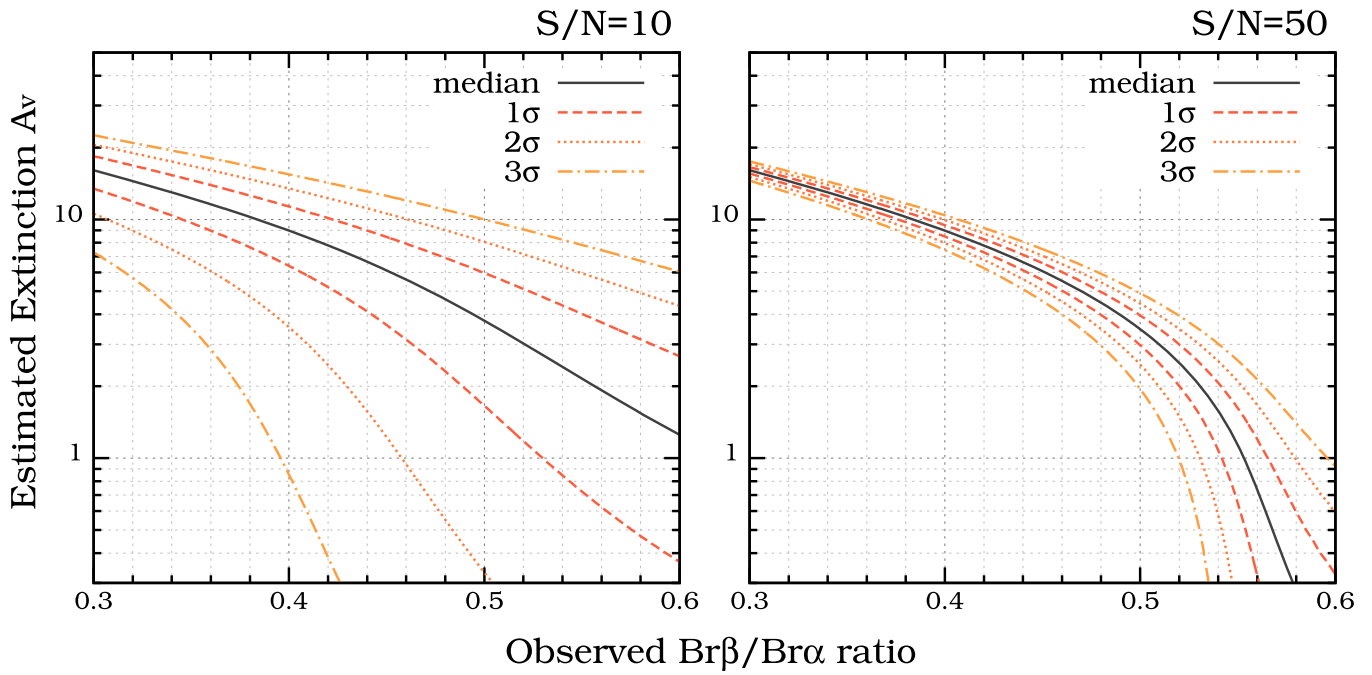
$$\begin{aligned} \text{prob}(\bar{x}_{\text{obs}} | x_{\text{obs}}, \sigma) &\propto \frac{1}{\sqrt{2\pi}\sigma} \\ &\times \exp\left(-\frac{(x_{\text{obs}} - \bar{x}_{\text{obs}})^2}{2\sigma^2}\right) \text{prob}(\bar{x}_{\text{obs}}), \end{aligned} \quad (14)$$

where  $\text{prob}(\bar{x}_{\text{obs}})$  is the prior distribution of  $\bar{x}_{\text{obs}}$  assumed to be a uniform distribution from 0 to  $x$ . From Equations (13) and (14), the posterior distribution of  $A_V(\text{Br}\beta, \text{Br}\alpha)$  is given by

$$\text{prob}(A_V | x_{\text{obs}}, \sigma) \propto \text{prob}(\bar{x}_{\text{obs}} | x_{\text{obs}}) \text{prob}(\bar{x}_{\text{obs}}). \quad (15)$$

The results for S/N = 10 and 50 are shown in Figure 19. The horizontal axis shows the observed intensity ratio ( $x_{\text{obs}}$ ), while the vertical axis is the estimated  $A_V$  value. The gray solid line indicates the median extinction  $A_V$ . The red dashed, orange dotted, and yellow dotted-dashed contours show the 1-, 2-, and 3 $\sigma$  confidence intervals, respectively. The typical S/N for the PNSPC spectrum is about 50. Figure 19 suggests that the extinction estimated from the Br $\alpha$ /Br $\beta$  ratio are quite uncertain when  $A_V \sim 1$ –3 mag. Figure 13 shows that the  $A_V$  values are typically less than about 2 mag for the PNSPC samples. Thus, it is practically impossible to precisely estimate the extinction of the PNSPC samples without using ancillary data such as the intensity of the H $\beta$  emission. The standard errors on  $A_V$  derived from the Br $\beta$ /Br $\alpha$  and H $\beta$ /Br $\alpha$  ratios are estimated to be about 0.5 and 0.02 mag., respectively. The intrinsic line ratio may depend on the electron temperature and density. The impact is much larger for the H $\beta$ /Br $\alpha$  ratio, but the variations in them can change  $A_V$  at most by 0.3 mag. Even though these variations are taken into account, the extinction derived from the H $\beta$ /Br $\alpha$  ratio is much more reliable than that from the Br $\beta$ /Br $\alpha$  ratio.





**Figure 19.** Estimated extinction and uncertainty. The dashed, dotted, and dotted–dashed contours show the confidence intervals of the  $1\sigma$ ,  $2\sigma$ , and  $3\sigma$ , respectively.

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